

YOUNG “DIPPER” STARS IN UPPER SCO AND ρ OPH OBSERVED BY K2M. ANSDELL^{1,10}, E. GAIDOS^{2,3,4,10}, S. A. RAPPAPORT⁵, T. L. JACOBS⁶, D. M. LACOURSE⁶, K. J. JEK⁶, A. W. MANN^{7,10}, M. C. WYATT⁸, G. KENNEDY⁸, J. P. WILLIAMS¹, T. S. BOYAJIAN⁹*Accepted to ApJ: October 28, 2015*

ABSTRACT

We present ten young ($\lesssim 10$ Myr) late-K and M dwarf stars observed in K2 Campaign 2 that host protoplanetary disks and exhibit quasi-periodic or aperiodic dimming events. Their optical light curves show ~ 10 – 20 dips in flux over the 80-day observing campaign with durations of ~ 0.5 – 2 days and depths of up to $\sim 40\%$. These stars are all members of the ρ Ophiuchus (~ 1 Myr) or Upper Scorpius (~ 10 Myr) star-forming regions. To investigate the nature of these “dippers” we obtained: optical and near-infrared spectra to determine stellar properties and identify accretion signatures; adaptive optics imaging to search for close companions that could cause optical variations and/or influence disk evolution; and millimeter-wavelength observations to constrain disk dust and gas masses. The spectra reveal Li I absorption and H α emission consistent with stellar youth (< 50 Myr), but also accretion rates spanning those of classical and weak-line T Tauri stars. Infrared excesses are consistent with protoplanetary disks extending to within ~ 10 stellar radii in most cases; however, the sub-mm observations imply disk masses that are an order of magnitude below those of typical protoplanetary disks. We find a positive correlation between dip depth and WISE-2 excess, which we interpret as evidence that the dipper phenomenon is related to occulting structures in the inner disk, although this is difficult to reconcile with the weakly accreting aperiodic dippers. We consider three mechanisms to explain the dipper phenomenon: inner disk warps near the co-rotation radius related to accretion; vortices at the inner disk edge produced by the Rossby Wave Instability; and clumps of circumstellar material related to planetesimal formation.

1. INTRODUCTION

Young stars exhibit diverse photometric variability attributed to a range of physical mechanisms. The bursting events seen in FU Orionis stars are likely due to episodic accretion during the early stages of stellar evolution (Hartmann & Kenyon 1996). The assortment of photometric variability observed in classical T Tauri stars (CTTS), whose light curves are complicated by their full disks and ongoing accretion, has been connected to a plethora of mechanisms, such as obscuring inner disk structures and accretion bursts (e.g., Cody et al. 2014). The sinusoidal light curves exhibited by more evolved weak-line T Tauri stars (WTTS), which no longer show strong accretion signatures, are attributed to magnetic spots on the stellar surface rotating with the star. Thus studying photometric variability in young stars can give insight into the structure and evolution of planet-forming

disks as well as the stellar activity affecting planets during their early evolution.

Studies of photometric variability in young stellar clusters have utilized space-based optical and infrared (IR) observatories, in particular the Convection, Rotation and Planetary Transits satellite (CoRoT; Baglin et al. 2006) and the *Spitzer* Infrared Array Camera (IRAC; Fazio et al. 2004), which provide more precise photometry and longer observing baselines than ground-based observatories. These studies have revealed a distinct class of photometric variables among CTTS populations, the so-called “dippers” whose light curves exhibit a relatively flat flux continuum but also contain quasi-periodic or aperiodic drops in flux lasting up to a few days.

In particular, the Orion Nebula Cluster (ONC) was observed by Morales-Calderón et al. (2011) with *Spitzer*/IRAC over 40 days and the NGC 2264 region was observed by Alencar et al. (2010) with CoRoT over 23 days and also by Cody et al. (2014) with *Spitzer*/IRAC and CoRoT simultaneously over 30 days. Morales-Calderón et al. (2011) found that the dips were shallower in the IR compared to the optical, consistent with extinction by dust; thus they interpreted the dips as arising from dusty inner disk structures passing along our line-of-sight (LOS) to the stellar photosphere. Alencar et al. (2010) and Cody et al. (2014) found such dippers to be fairly common (~ 20 – 30%) among the CTTS population in NGC 2264. Stauffer et al. (2015) and McGinnis et al. (2015) later attributed the dippers in NGC 2264 to occultations of the stellar photosphere by disk structures at or near the co-rotation radius.

Detailed follow-up investigations of the dipper populations in young star-forming regions are needed to improve

¹ Institute for Astronomy, University of Hawai‘i at Mānoa, Honolulu, HI

² Department of Geology & Geophysics, University of Hawai‘i at Mānoa, Honolulu, HI

³ Visiting Scientist, Sauverny Observatoire, l’Université de Genève

⁴ Visiting Scientist, Center for Astrophysics, Harvard University, Cambridge, MA

⁵ Physics Department and Kavli Institute for Astrophysics and Space Research, Massachusetts Institute of Technology, Cambridge, MA

⁶ Amateur Astronomer

⁷ Harlan J. Smith Fellow, Department of Astronomy, The University of Texas at Austin, Austin, TX

⁸ Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge, UK

⁹ Department of Astronomy, Yale University, New Haven, CT

¹⁰ Visiting Astronomer, NASA Infrared Telescope Facility, operated by the University of Hawaii under contract NNN14CK55B

Table 1
Properties of K2/C2 Dippers

EPIC ^a	2MASS ^b	RA _{J2000}	Dec _{J2000}	Mem. ^c	P _{rot} ^d	N _{dip} ^e	R _{dip} ^e	D _{dip} ^e	Var. ^f
203343161	16245587-2627181	16:24:55.878	-26:27:18.10	USc	2.24	14	10.2	0.12	Q
203410665	16253849-2613540	16:25:38.484	-26:13:53.98	Oph	4.24	17	7.7	0.30	A
203895983	16041893-2430392	16:04:18.936	-24:30:39.34	USc	2.44	10	5.5	0.11	A
203937317	16261706-2420216	16:26:17.080	-24:20:22.02	Oph	5.44	17	11.8	0.13	Q
204137184	16020517-2331070	16:02:05.180	-23:31:06.94	USc	2.64	20	9.4	0.35	Q
204630363	16100501-2132318	16:10:05.019	-21:32:31.89	USc	6.66	8	7.8	0.10	A
204757338	16072747-2059442	16:07:27.462	-20:59:44.22	USc	2.39	12	7.1	0.08	Q
204932990	16115091-2012098	16:11:50.928	-20:12:09.89	USc	2.30	11	15.3	0.26	A
205151387	16090075-1908526	16:09:00.762	-19:08:52.70	USc	9.55	15	8.8	0.31	Q
205519771	16071403-1702425	16:07:14.018	-17:02:42.67	USc	2.46	13	11.5	0.11	A

^a K2 Ecliptic Plane Input Catalog (EPIC) ID.

^b 2MASS ID associated with the EPIC ID.

^c Cluster membership: USc = Upper Sco, Oph = ρ Oph (Section 2.4).

^d Inferred rotation period (days) measured from the K2/C2 light curves (Section 2.2).

^e Criteria used to select our dipper sample (Section 2.2).

^f Variability type: A = aperiodic, Q = quasi-periodic (Section 2.3)

our physical understanding of the dipper phenomenon, which in turn can provide a means of probing disk structure and dynamics during key epochs of planet formation. For example, if the dips are indeed produced by occulting inner disk structures, they could provide a rare opportunity to study the inner disk and its role (if any) in planet formation. If the dips are instead related to material further out in the disk, they could provide information on the size distribution of forming planetesimals in debris disks or transiently formed vortices near the inner edges of transition disks.

However, the ONC and NGC 2264 are both young and distant clusters. The ONC has an age of ~ 2 Myr and a distance of ~ 400 pc (e.g., [Reggiani et al. 2011](#)), while NGC 2264 has an age of ~ 3 Myr and a distance of ~ 760 pc (e.g., [Dahm 2008](#)). Their distances limit detailed follow-up observations of their dipper populations (e.g., resolving ~ 30 AU scales in NGC 2264 to search for close companions requires ~ 40 mas resolution), while their similar ages make it desirable to observe other star-forming regions at distinctly different ages (e.g., to study the evolution of the dipper phenomenon).

Fortunately, the K2 mission ([Howell et al. 2014](#)) provides a new opportunity for studying photometric variability in young stellar clusters. Although the original mission of the *Kepler* spacecraft ([Borucki et al. 2010](#)) focused on solar-age stars at several hundred parsecs, the re-purposed K2 mission is studying fields that include nearby young stellar clusters with ages spanning ~ 1 –600 Myr. In particular, the K2 Campaign 2 (K2/C2) field covered the Upper Scorpius (Upper Sco) and ρ Ophiuchus (ρ Oph) regions, which are relatively nearby at ~ 140 pc ([de Zeeuw et al. 1999](#)) and ~ 120 pc ([Loinard et al. 2008](#)), respectively, and have significantly different ages at ~ 10 Myr ([Pecaut et al. 2012](#)) and ~ 1 Myr ([Andrews & Williams 2007](#)), respectively.

This work focuses on a sample of dippers in the Upper Sco and ρ Oph regions found using the publicly available K2/C2 data. We describe our dipper sample selection in Section 2, then present our follow-up observations in Section 3. We analyze these data in Section 4 and identify observed trends and correlations in Section 5. In Section 6, we discuss possible mechanisms driving the dipper phenomenon. We summarize our findings and

discuss areas of future work in Section 7.

2. DIPPER SAMPLE

2.1. K2 Light Curve Extraction

The primary mission of the *Kepler* spacecraft ended in 2013 May, when the failure of two of its four reaction wheels meant the spacecraft could no longer maintain the high photometric precision required for its original science goals. The re-purposed two-wheeled *Kepler* mission, K2, utilizes an ecliptic-observing orientation to mitigate pointing drift due to solar radiation pressure. However, quasi-periodic thruster firings are still needed to correct for residual pointing drift throughout the 80-day observing campaign. This spacecraft motion introduces artificial systematics to the K2 light curves as differing amounts of target flux may be lost/contaminated when applying aperture masks to the K2 Target Pixel Files (TPFs). Although artifact mitigation techniques were previously developed for the *Kepler* mission (e.g., [Kinnouchi et al. 2012](#)), these additional pointing variations require new data reduction solutions optimized for K2 (e.g., [Vanderburg & Johnson 2014](#); [Huang et al. 2015](#)).

We therefore employed a fixed-mask data reduction pipeline based on the open source *Kepler* Guest Observer software package PyKE¹¹ ([Still & Barclay 2012](#)), which was updated in 2014 March to provide additional compatibility with K2-derived TPFs. We obtained K2/C2 data from the Mikulski Archive for Space Telescopes¹² (MAST) in 2015 March, and utilized the Principal Component Analysis (PCA) technique KEPPCA¹³ ([Harri-son et al. 2012](#)) to extract light curves for each of the 13,344 stellar long cadence TPFs. KEPPCA mitigates artifacts by analyzing pixel samples from a given TPF and segregating observed photometric trends associated with individual pixels from those common to all pixels. We extracted source fluxes using square aperture masks centered at the source locations, where mask dimensions were adjusted as a function of EPIC catalog source magnitude. Our pipeline automatically identifies additional

¹¹ <http://keplerscience.arc.nasa.gov/PyKE.shtml>

¹² <https://archive.stsci.edu/>

¹³ KepPCA release notes: <http://keplerscience.arc.nasa.gov/ContributedSoftwareKepPCA.shtml>

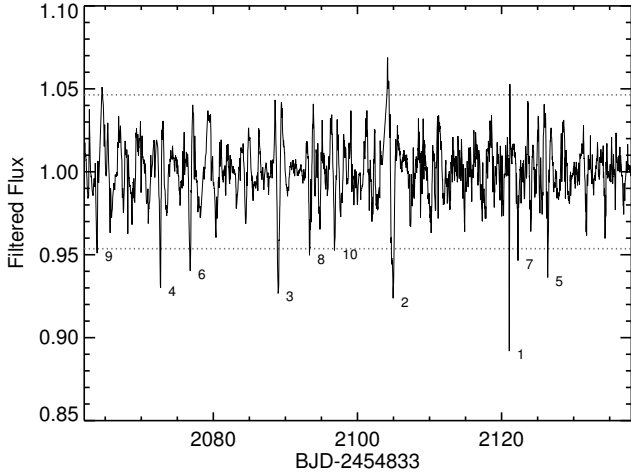


Figure 1. Filtered light curve for EPIC 203895983. The x axis shows the Barycentric Julian Date (BJD) spanning the 80-day K2/C2 observing period. The y axis shows the normalized flux after applying a high-pass filter (Section 2.2). The dips in flux below the 3σ threshold (bottom dotted line) are numbered in order of dip depth. For comparison, the unfiltered light curve for EPIC 203895983 is shown in Figure 2.

centroids in the mask, but does not subsequently adjust the mask to avoid contamination from nearby sources before extracting photometry.

With any de-trending process comes the risk of altering the underlying astrophysical signals. Thus we compared our corrected data to the K2/C2 K2SFF¹⁴ light curves released to MAST in 2015 May, which have been motion corrected with the Self Field Flattening (SFF) technique described in Vanderburg & Johnson (2014). The SFF method improves the pointing precision of K2 aperture photometry by correlating observed variability with spacecraft motion, and has revealed subtle long-term drifts in target centroid positions not accounted for in our KEPPCA pipeline. Therefore we used the KEPPCA light curves when identifying our dipper sample (Section 2.2), but used the K2SFF light curves for all subsequent analyses; the exception is EPIC 204137184, whose K2SFF light curve appeared corrupted.

2.2. Dipper Sample Identification

We reviewed each of the 13,344 corrected light curves (Section 2.1) by eye, flagging ~ 100 sources exhibiting quasi-periodic or aperiodic dimming events as candidate dippers. The visual inspection was necessary because dippers exhibit complex light curves that make them difficult to identify with automated algorithms. We then conservatively ignored light curves showing suspicious irregularities that appeared to be instrument related (e.g., charge bleed) or the result of data corruption (e.g., data discontinuities) rather than due to any physical phenomenon inherent to the science target. We also ignored particularly noisy light curves that were likely associated with extreme intrinsic stellar variability. For each candidate dipper, we normalized the extracted K2/C2 light curve, then put the normalized light curve through a high-pass filter with a cut-on frequency of 1 day^{-1} . This filtering highlighted the quasi-periodic and aperi-

odic dimming events, while suppressing the periodic variability from stellar rotation, due to their different duty cycles. Figure 1 shows an example of a filtered light curve (to be compared with the normalized, but unfiltered, versions in Figure 2). We used these light curves to compute several metrics that quantify the dipper phenomenon, and then used the metrics to select our sample of ten dippers presented in Table 1. We describe these metrics and our selection criteria below.

To estimate the inherent variation of each source (σ) we calculated the Tukey’s biweight function (Tukey 1977) of the filtered light curve using the IDL function `robust_sigma`, which gives a weighted standard deviation. The number of dips over the 80-day observing period (N_{dip}) was tallied from the drops in flux at least 3σ in depth in the filtered light curve (see Figure 1). We calculated the average depth of the three deepest dips in the normalized light curve (D_{dip}), and the ratio of the average depth of the three deepest dips in the filtered light curve to σ (R_{dip}), as measures of the strength of the dipper phenomenon. We also estimated the stellar rotation period (P_{rot}) by computing the autocorrelation function of the normalized light curve, checking our results with Fourier Transform and Lomb-Scargle periodogram techniques to ensure consistency. This assumes that the periodic signal is from stellar rotation and is not influenced by the dimming events; thus although our P_{rot} values are consistent with the stellar rotation periods of young late-type stars (e.g., Herbst et al. 2007), they should be viewed with caution until confirmed by other observations (e.g., $v \sin i$ from high-resolution spectra).

To construct our dipper sample, we required $N_{\text{dip}} > 5$ (to ensure dips were not spurious events), $R_{\text{dip}} > 5$ and $D_{\text{dip}} > 0.07$ (to identify significant dippers), and $P_{\text{rot}} \lesssim 10$ days (to confirm stellar youth; Herbst et al. 2007). This resulted in a sample of ten dippers, whose properties are listed in Table 1 and whose normalized light curves are shown in Figure 2. We checked for planetary signals using the IPAC Box-Least Squares (BLS) periodogram service¹⁵ but did not find any planetary signals.

2.3. Dipper Classification

The K2/C2 light curves of the dippers in our sample (Figure 2) show quasi-periodic or aperiodic dimming events that appear as discrete dips in flux with typical durations of ~ 0.5 – 2.0 days and amplitudes of up to $\sim 40\%$. Following Cody et al. (2014) and McGinnis et al. (2015), we define quasi-periodic dippers as those with dimming events that appear at periodic intervals but with varying shapes and depths, and aperiodic dippers as those with dimming events that appear stochastically and with varying shapes and depths. To classify the periodicity, we fit a hyperbolic secant function with time-varying dip depths to each light curve:

$$y = F_o - \sum_n 2D(t_n) \times \left[\exp\left(\frac{t - t_n}{\tau}\right) + \exp\left(-\frac{t - t_n}{\tau}\right) \right]^{-1}, \quad (1)$$

where $t_n = t_o + nP_{\text{rot}}$. Here F_o is the baseline flux

¹⁴ <https://archive.stsci.edu/prepds/k2sff/>

¹⁵ <http://exoplanetarchive.ipac.caltech.edu/applications/Periodogram/>

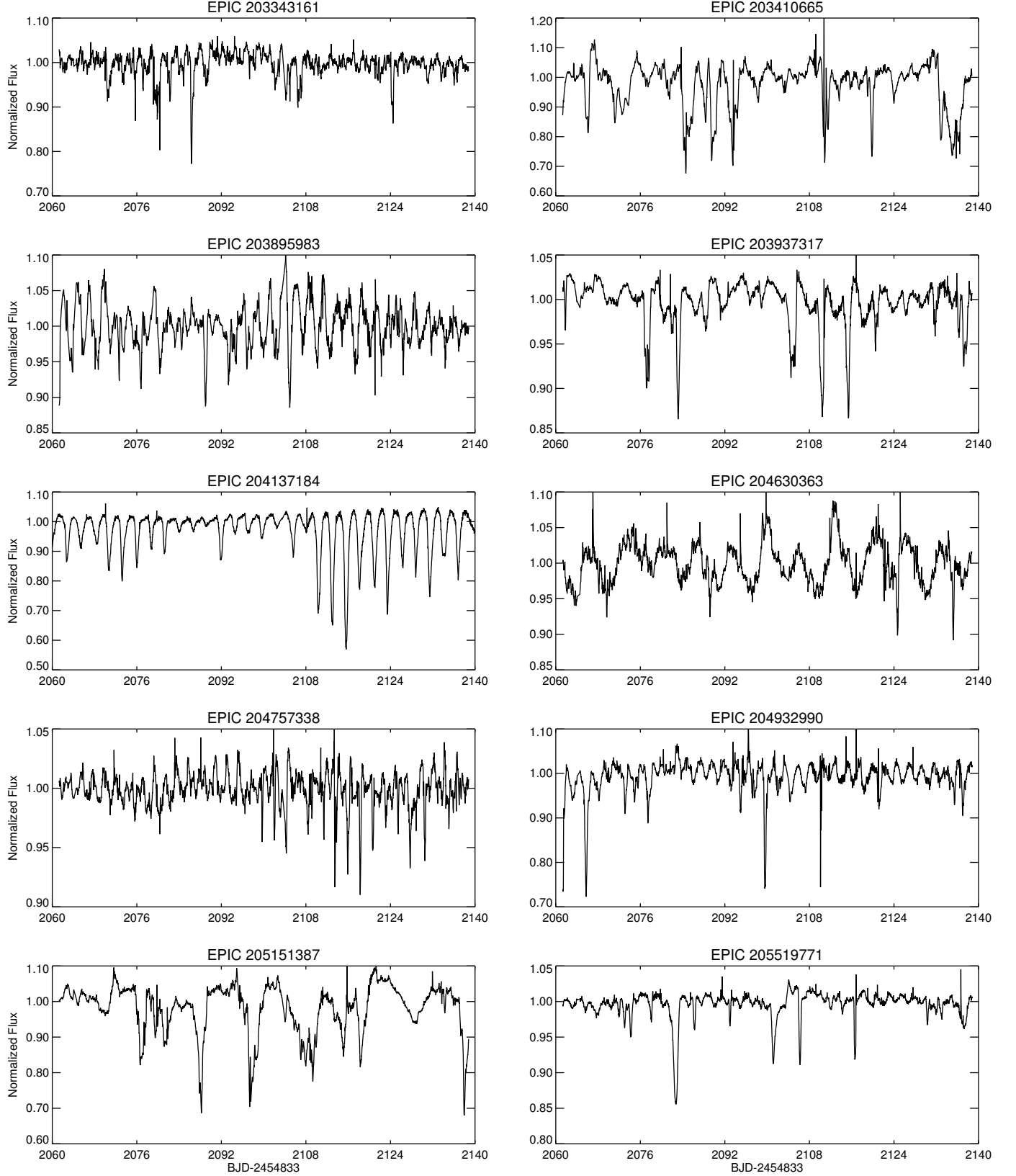


Figure 2. Normalized light curves (Section 2.2) for the ten dippers in our sample (Table 1). Photometry is from the publicly available K2SFF light curves, except for EPIC 204137184 whose photometry is from our KEPKA pipeline (Section 2.1). The narrow vertical spikes seen in some dippers are likely cosmic rays, data processing artifacts, or stellar flares.

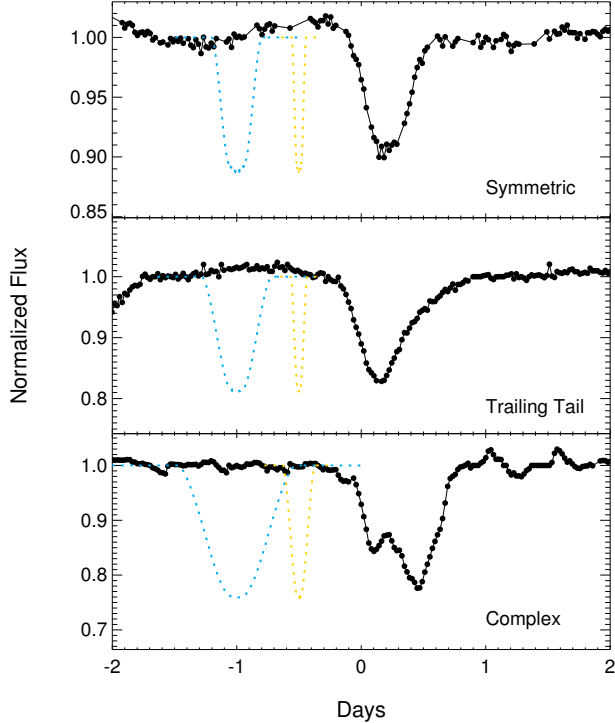


Figure 3. Examples of the three main dip types seen in the K2/C2 light curves (black points). We also plot model planetary transits (offset for clarity) of the same depth to show their inconsistencies with the dips (see also Section 6). Yellow transit curves have durations based on circular orbits with periods equal to our P_{rot} value for the source, while blue transit curves have durations extended to match the FWHM of the dip. Model transits use the Claret (2004) limb-darkening laws interpolated for the T_{eff} value of the star with an impact parameter of 0.5. Symmetric dips (top) resemble planet transits in shape, but require transit durations that are too long for the orbital period as well as transit depths that are too deep to be planetary. Asymmetric dips with lagging tails (middle) or complex structures (bottom) are clearly non-planetary.

level, t_o is the location of the first dip in the light curve, n is the dip number, P_{rot} is the inferred rotation period, τ is the dip duration, and $D(t_n)$ is the depth of dip n . We used the MPFIT implementation of the Levenberg-Marquardt technique for χ^2 minimization (Markwardt 2009) to find the best-fit parameters for each source while fixing P_{rot} to our derived value (Section 2.2; Table 1).

As shown in Figure 4, this function can reasonably reproduce the dip occurrence, width, and depth for quasi-periodic but not aperiodic dippers. We found that some dippers (e.g., EPIC 203937317) have dips that appear episodic in their K2/C2 light curves, but then occur in phase when the light curve is phase folded to their P_{rot} value; we consider these dippers as quasi-periodic, as this implies that the dimming events occur at periodic intervals (though not at every interval). Table 1 gives the periodicity classification (quasi-periodic or aperiodic) for each dipper in our sample.

Figure 3 shows examples of the three main dip types seen in the K2/C2 light curves, illustrating that the dips are clearly inconsistent with planetary transits.

2.4. Cluster Membership

The dippers in our sample are all members of either the Upper Sco or ρ Oph star-forming regions (Table 1), despite being selected only on their optical light curve properties. EPIC 203410665 (V* V896 Sco) and 203937317 (DoAr 24) are both well-known members of ρ Oph that have been studied in the near-IR (e.g., Furlan et al. 2009; McClure et al. 2010). EPIC 205151387 is a well-studied member of Upper Sco whose disk has been detected with *Herschel* PACS (Mathews et al. 2013), the Sub-millimeter Array (SMA; Cieza et al. 2008), and the Atacama Large Millimeter Array (ALMA; Carpenter et al. 2014). EPIC 203895983, 204630363, and 205519771 were spectroscopically confirmed as low-mass members of Upper Sco by Rizzuto et al. (2015). EPIC 204757338 and 204932990 were part of the Luhman & Mamajek (2012) sample of disk-bearing members in Upper Sco. EPIC 203343161 was classified as an Upper Sco member based on its photometry and proper motion (Lodieu 2013). EPIC 204137184 was identified as a candidate low-mass member of Upper Sco, though based only on photometry (Ardila et al. 2000).

3. FOLLOW-UP OBSERVATIONS

We obtained multi-wavelength follow-up observations of our dippers, namely: optical and near-IR spectra (Section 3.1) to determine stellar properties and accretion states; sub-millimeter observations (Section 3.2) to constrain disk dust and gas masses; and adaptive optics (AO) imaging (Section 3.3) to search for close companions that could influence optical variations and/or disk evolution. We also utilized various literature and public data sources (Section 3.4). These follow-up observations are described below and summarized in Table 2.

3.1. Spectroscopy

We obtained optical spectra for our complete dipper sample using the Super-Nova Integral Field Spectrograph (SNIFS; Aldering et al. 2002; Lantz et al. 2004) at the University of Hawaii 2.2 m telescope atop Maunakea during 2015 May–July. These spectra have resolutions $R \approx 900$ and signal-to-noise ratios $\text{SNR} \gtrsim 30$ per resolution element at $\gtrsim 5100$ Å. Details of our SNIFS data reduction can be found in Mann et al. (2012) and Lépine et al. (2013). We used these spectra to estimate stellar properties (Section 4.1) and identify accretion states (Section 4.2). We performed spectral monitoring over several months (Table 2) to study activity variability and account for changes in apparent effective temperature (T_{eff}) due to variations in star spot coverage.

We used the Echellette Spectrograph and Imager (ESI) at the Keck II 10 m telescope atop Maunakea to obtain optical spectra of seven dippers in our sample on 2015 July 17. We used the instrument in cross-dispersed mode with a $0.5''$ slit to achieve $R \approx 8000$. Integration times of a few minutes yielded $\text{SNR} \sim 20$ –90 in the Li I 6708 Å line (Table 2). Spectra were reduced using the white dwarf standard EG 131 with the publicly available ESIRedux pipeline¹⁶ with dome flats for flat-fielding and arc lamps for wavelength calibration (Prochaska et al.

¹⁶ <http://www2.keck.hawaii.edu/inst/esi/ESIRedux/index.html>

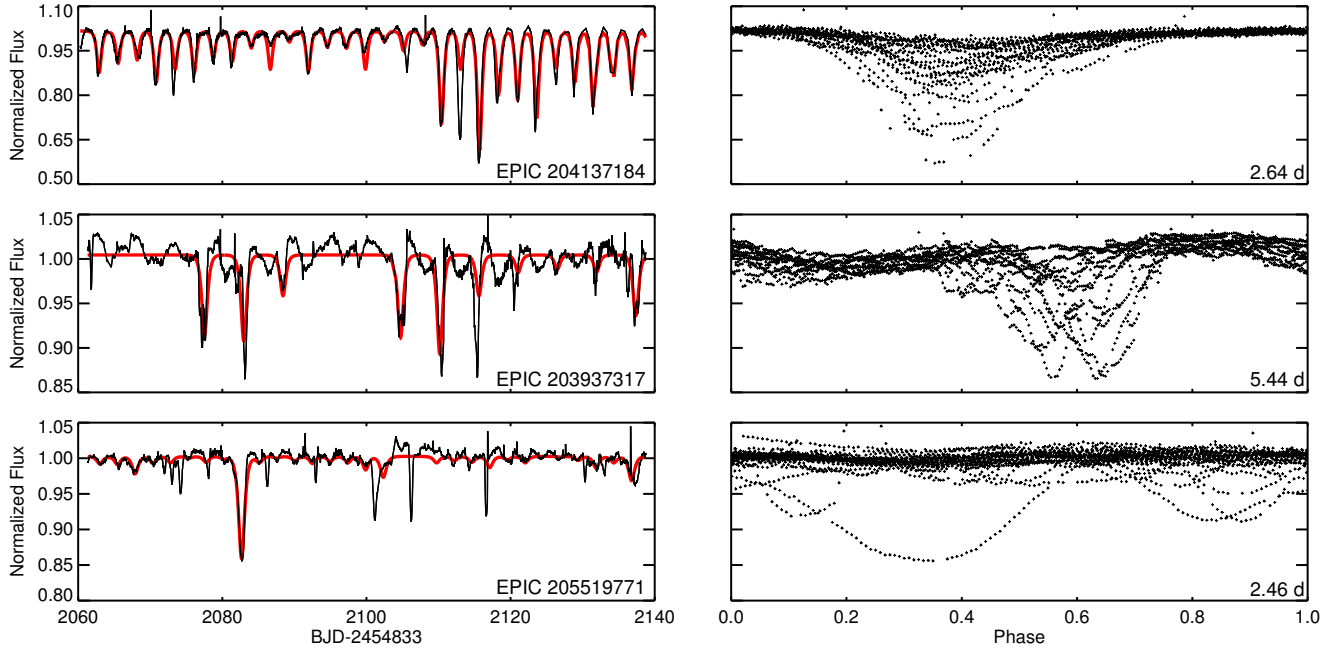


Figure 4. Examples of quasi-periodic (top, middle) and aperiodic (bottom) dippers. Red lines show fits to the normalized light curves (left) using a hyperbolic secant function with time-varying dip depths (Equation 1; Section 2.3). The function can reasonably reproduce dip occurrence, width, and depth for quasi-periodic but not aperiodic dippers. Phase-folded light curves (right) show dips occurring at similar phases for quasi-periodic but not aperiodic dippers. The P_{rot} values used for phase folding are given for reference (see also Table 1 & Section 2.2).

2003; Bochanski et al. 2009). Our ESI spectra were used to measure Li I equivalent widths (Section 4.1) and search for spectroscopic binaries (Section 4.6).

We acquired near-IR spectra for eight dippers in our sample during 2015 May–July using the upgraded SpeX spectrograph (Rayner et al. 2003) on the 3 m NASA Infrared Telescope Facility (IRTF) atop Maunakea. uSpeX spectra were taken in short cross-dispersed (SXD) mode using the $0.3'' \times 15''$ slit, covering 0.7 to 2.5 μm at $R \approx 2000$. These spectra have $\text{SNR} \gtrsim 40$ in K band (Table 2). Basic reduction (bias, flat correction, extraction, etc.) was carried out with SpeXTool (Cushing et al. 2004). Flux calibration and telluric line removal were performed using A0V standards with Xtelcor (Vacca et al. 2003). Details on the observation, reduction, and extraction methods can be found in Mann et al. (2015). We primarily used our uSpeX spectra to check for near-IR accretion signatures (Section 4.2).

3.2. Sub-millimeter Observations

We obtained sub-mm observations with the SMA (Ho et al. 2004), an interferometer on Maunakea consisting of eight 6 m antennas. Data were obtained during 2015 June–July for all dippers in our sample except EPIC 205151387, which was previously observed with the SMA by Cieza et al. (2008). We used the SMA 1.3 mm band with the correlator configured to cover CO 2–1 line emission at 230.538 GHz. We used sub-compact or compact array configurations with baselines of ~ 25 m or ~ 50 m, respectively. The precipitable water vapor was typically ~ 2.5 mm, corresponding to an atmospheric optical depth at zenith of ~ 0.1 at 225 GHz. Each source was observed for 6 min in a cycle that was interleaved with 3 min integrations on quasars 1517-243 and 1625-254 for calibrating the variation of instrumental amplitude and phase

with time. The tracks spanned hour angles of roughly -2.5^{h} to $+2.5^{\text{h}}$ resulting in total integration times of ~ 1 hr per source. The bandpasses were calibrated at the end of the tracks with 30 min integrations on 3C454.3 and the absolute flux scales were derived from observations of Titan.

We used the MIR software package¹⁷ to calibrate the data and the MIRIAD software package¹⁸ for imaging. The close configuration of the array resulted in large beam sizes of roughly $8'' \times 5''$. The achieved rms noise level for each source is given in Table 2. We used our SMA observations to constrain disk dust and gas masses (Section 4.3), which are important for classifying disk evolutionary states.

3.3. Adaptive Optics Imaging

We obtained AO imagery for five dippers in our sample with the Near-infrared Imaging Camera (NIRC2) on the Keck II 10 m telescope atop Maunakea on 2015 June 22 using laser guide star AO (LGS-AO; Wizinowich et al. 2006; van Dam et al. 2006). Imaging was done in K' band with the narrow camera, which has a pixel scale of 9.952 ± 0.002 mas pixel $^{-1}$ and an orientation angle of $0^\circ 252 \pm 0^\circ 009$ (Yelda et al. 2010). Weather conditions were photometric with seeing $< 0.5''$. These NIRC2 images can reveal objects with $\Delta K' \sim 7\text{--}8$ mags at 1000 mas or $\Delta K' \sim 4$ mags at 100 mas. At the distance of Upper Sco and ρ Oph, ~ 100 mas corresponds to ~ 15 AU, thus we used these data to search for close companions that could contribute to optical variations and/or influence disk evolution (Section 4.6).

¹⁷ <https://www.cfa.harvard.edu/~cqi/mircook.html>

¹⁸ <https://www.cfa.harvard.edu/sma/miriad>

3.4. Archival Data

All dippers in our sample were observed by the Two Micron All-Sky Survey (2MASS; [Skrutskie et al. 2006](#)), supplying near-IR photometry in JHK_S bands. The dippers were also detected by the Wide-field Infrared Survey Explorer (WISE; [Wright et al. 2010](#)), providing mid-IR photometry at 3.4, 4.6, 12, and 22 μm bands, denoted as W1 through W4. EPIC 203410665, 203937317, and 205151387 were observed with *Spitzer* IRS, providing mid-IR spectra covering the silicate emission feature at 10 μm . EPIC 205151387 was observed with *Herschel* PACS at 70, 100, and 160 μm bands ([Mathews et al. 2013](#)). For EPIC 203937317, we were able to estimate its PACS 70 μm flux from a *Herschel* archive image. EPIC 205151387 was detected by ALMA at 880 μm ([Carpenter et al. 2014](#)) and the SMA at 1.3 mm ([Cieza et al. 2008](#)). These archival data are summarized in Table 6.

4. ANALYSIS

4.1. Stellar Properties

We estimated T_{eff} values for each dipper following the procedure in [Mann et al. \(2013\)](#) and [Gaidos et al. \(2014\)](#). This method can reproduce interferometric T_{eff} measurements of nearby M dwarfs ([Boyajian et al. 2012](#)) to $\simeq 60$ K, but is untested on young ($\lesssim 10$ Myr) stars. In short, we compared our SNIFS spectra to the CFIST grid of the BT-SETTL version of the PHOENIX stellar atmosphere models ([Allard et al. 2013](#)) with the relative abundances of solar elements as estimated by [Caffau et al. \(2011\)](#). We explored a grid of $\log g$ from 3.5 to 5.0 and T_{eff} from 2500 to 5000 K, but restricted metallicity and alpha abundance to solar values because large deviations from solar are unlikely for nearby, pre-main sequence stars. We fit and divided out a third-order polynomial in flux to capture the effects of reddening and imperfect flux calibration as well as a linear term in wavelength to correct for small errors in wavelength calibration and radial velocity offsets. The T_{eff} value for a given spectrum was taken from the best-fit model, with measurement errors typically $\lesssim 100$ K. However, as a star rotates we observe different star spot distributions, which can change the derived T_{eff} value due to the cooler temperature of star spots relative to the surrounding photosphere. Thus for each dipper we obtained T_{eff} values for each SNIFS observation epoch (Table 2), taking the weighted mean of these values as the final T_{eff} estimate reported in Table 3. Errors were determined by adding in quadrature the standard error on the mean and the aforementioned typical measurement errors.

Spectral subtypes were determined from spectral index measurements of molecular bands in our SNIFS spectra. Specifically, we applied correlations between spectral subtype and the depths of CaH and TiO molecular bands derived in [Lépine et al. \(2013\)](#). Reddening and spectral subtype are highly degenerate when using only photometry, however the application of spectral index measurements is much more robust because changes in reddening cannot reproduce the depth of the molecular bands. Although spectral indices can be affected by the lower gravity of young objects, our estimated spectral subtypes reported in Table 3 are consistent with our derived T_{eff} values and accurate to \pm one spectral subtype.

We estimated reddening (A_V) by comparing our SNIFS

Table 2
Follow-up Observations

EPIC ID	SNIFS ^a	ESI ^b	uSpeX ^c	SMA ^d	NIRC2 ^e
203343161	4	25	40	1.9	N
203410665	7	—	250	1.6	Y
203895983	7	—	185	1.5	Y
203937317	6	90	75	1.8	Y
204137184	5	30	—	1.5	N
204630363	5	—	—	1.5	N
204757338	5	20	110	1.4	N
204932990	7	30	100	1.6	Y
205151387	5	60	125	2.0	N
205519771	7	30	100	1.3	Y

^a Number of SNIFS observation epochs (Section 3.1).

^b SNR of ESI spectra in the Li I 6708 Å line (Section 3.1)

^c SNR of uSpeX spectra in K band (Section 3.1).

^d Flux rms noise (mJy) of SMA observations (Sections 3.2 & 3.4)

^e NIRC2 imagery (Section 3.3): Y = observed, N = not observed.

spectra to a sequence of template M dwarf spectra from the nearby TW Hydrae Association (TWA) taken as part of a spectroscopic study of nearby young moving groups (Mann et al., in prep). The TWA members are within the local bubble where interstellar extinction is negligible ([Aumer & Binney 2009](#)) and have ages similar to the dippers in our sample (~ 10 Myr; [Weinberger et al. 2013](#); [Ducourant et al. 2014](#)). Template spectra were taken with the same instrument and reduced with the same methods as the dipper spectra. We masked out regions with strong emission (e.g., H α at 6563 Å) or telluric absorption features (e.g., O $_2$ near 7600 Å), then artificially reddened each template spectrum using an $R_V = 3.1$ reddening law from [Cardelli et al. \(1989\)](#). We marginalized over the template choice, A_V , and radial velocity. The reddened templates reasonably reproduced the dipper spectra with typical best-fit reduced χ^2 of ~ 1.6 , however the strong degeneracy between template spectral subtype and A_V resulted in larger errors. Figure 5 illustrates the fitting procedure and final estimates on A_V are given in Table 3.

We measured equivalent widths of the Li I line at 6708.0 Å (EW_{Li}) from our higher-resolution ESI spectra. The presence of Li I absorption is an indicator of stellar youth (< 50 Myr) in late-type stars because the element is rapidly convected into the stellar core where it is destroyed. We simultaneously fit two Gaussians, one centered on the Li I line at 6708.0 Å (our ESI resolution was too low to resolve the doublet) and the other centered on the nearby Fe I line at 6707.5 Å. We used the fitted Li I line to estimate EW_{Li} values while accounting for any Fe I contamination. Errors were determined using a Monte Carlo method: we used the standard deviation of the flux in two continuum regions flanking the lines to add Gaussian-distributed noise to the observed spectrum, then repeated the fitting procedure 100 times, taking the mean and standard deviation as our final EW_{Li} value and error, respectively. The EW_{Li} values reported in Table 3 are consistent with those found for M dwarfs in Upper Sco (e.g., see Figure 5 in [Rizzuto et al. 2015](#)).

4.2. Accretion Indicators

Spectral accretion signatures are produced by shocked gas falling onto the star or encountering the stellar magnetosphere, making them useful indicators of inner

Table 3
Stellar and Disk Properties

EPIC	$T_{\text{eff}}^{\text{a}}$ (K)	SpT ^b	A_V^{c} (mag)	$\text{EW}_{\text{H}\alpha}^{\text{d}}$ (Å)	$\text{EW}_{\text{Li}}^{\text{e}}$ (Å)	He I ^f	$\text{EW}_{\text{Pa}\gamma}^{\text{g}}$ (Å)	$M_{\text{dust}}^{\text{h}}$ (M_{\oplus})	Acc. ⁱ	Disk. ^j
203343161	3120±115	M5.5	0.09 ^{+0.13} _{-0.09}	23.4 ⁺¹⁰ _{-8.9}	-0.46±0.03	C	0.6±0.3	< 3	W/C	Ev
203410665	4045±50	K7.0	1.00 ^{+0.13} _{-0.35}	1.1 ^{+1.9} _{-1.4}	—	R	1.3±0.1	< 3	W	F/(P)TD
203895983	3655±75	M2.5	0.31 ^{+0.19} _{-0.06}	11.5 ^{+3.4} _{-6.1}	—	R/B	0.4±0.1	< 3	W/C	F/(P)TD
203937317	4070±50	K7.5	2.57 ^{+0.13} _{-0.16}	15.3 ^{+3.6} _{-3.2}	-0.39±0.02	C	1.2±0.2	4	C	F
204137184	3210±115	M4.5	0.41 ^{+0.09} _{-0.13}	2.7 ^{+0.8} _{-1.0}	-0.60±0.03	—	—	< 3	W	F
204630363	4070±50	K7.5	0.81 ^{+0.13} _{-0.13}	24.2 ⁺¹¹ _{-8.2}	—	—	—	15	C	F/(P)TD
204757338	3145±100	M4.5	1.03 ^{+0.09} _{-0.06}	4.3 ^{+1.3} _{-1.2}	-0.49±0.03	R	0.0±0.1	7	W	F/(P)TD
204932990	3390±115	M3.5	0.88 ^{+0.31} _{-0.09}	6.9 ^{+8.6} _{-3.7}	-0.40±0.04	R/B	0.0±0.1	< 3	W	F/(P)TD
205151387	3975±130	M1.0	0.78 ^{+0.13} _{-0.39}	6.9 ^{+2.4} _{-5.0}	-0.47±0.02	B	1.0±0.1	7–9	W	F/(P)TD
205519771	3390±100	M3.5	0.78 ^{+0.13} _{-0.34}	2.5 ^{+1.1} _{-0.8}	-0.48±0.04	R	0.1±0.1	< 3	W	Ev

^a Effective temperature measured from our SNIFS spectra (Section 4.1).

^b Spectral subtypes estimated from our SNIFS spectra, accurate to \pm one spectral subtype (Section 4.1).

^c Extinction estimated from our SNIFS spectra (Section 4.1; Figure 5).

^d Mean H α equivalent widths with high/low ranges observed across our SNIFS spectral monitoring campaign (Section 4.2; Table 2)

^e Li I equivalent widths measured from our ESI spectra (Section 4.1).

^f He I accretion signatures seen in our uSpeX data: C = centered absorption, R = redshifted absorption, B = blueshifted absorption (Section 4.2)

^g H I Pa γ equivalent widths measured from our uSpeX spectra (Section 4.2)

^h Disk dust masses estimated from our SMA data; for EPIC 205151387, we give the range found in the literature (Section 4.3).

ⁱ Accretion state based on $\text{EW}_{\text{H}\alpha}$ and spectral subtype (criterion from White & Basri 2003): W = WTTS, C = CTTS. We use W/C to indicate sources that showed a large range of $\text{EW}_{\text{H}\alpha}$ values during our spectral monitoring campaign and thus could be classified as WTTS or CTTS.

^j Disk type determined from IR color excesses (Section 4.4) and SEDs (Section 4.5): F = Full, TD = Transition Disk, Ev = Evolved. We use F/(P)TD to indicate disks that were classified as “full” by their IR colors but have SEDs indicative of (pre-) transition disks.

gaseous disks and stellar evolutionary states. We first searched our SNIFS spectra for H α emission from H I at 6563 Å, which is exhibited by all our dippers. However, certain levels of H α emission can be emitted by young, late-type stars that are active but whose disks have dissipated and thus are no longer accreting. White & Basri (2003) found that for accreting CTTS, H α equivalent widths ($\text{EW}_{\text{H}\alpha}$) were ≥ 10 Å for K7–M2.5, ≥ 20 Å for M3–M5.5, and ≥ 40 Å for M6 and later spectral subtypes. We therefore obtained $\text{EW}_{\text{H}\alpha}$ values from our SNIFS spectra following the procedure in Lépine et al. (2013): we measured flux over a 14 Å-wide region (6557.61–6571.61 Å) relative to pseudo-continuum regions (6500–6550, 6575–6625 Å) and calculated errors using the Monte Carlo method described in Section 4.1. Table 3 gives the average and high/low ranges of the $\text{EW}_{\text{H}\alpha}$ values for each dipper across all SNIFS observation epochs. These levels of H α emission are more consistent with stellar youth rather than ongoing accretion: according to the criteria of White & Basri (2003), only two dippers (EPIC 203937317, 204630363) consistently exhibited H α emission at levels expected for accreting CTTS.

We searched our uSpeX spectra for near-IR accretion signatures, focusing on Pa γ emission from H I at 1.094 μm and He I absorption at 1.083 μm . Pa γ emission arises from magnetospheric accretion as with H α . He I absorption at 1.083 μm is thought to be particularly sensitive to inner disk flows due to the metastability of the transition’s lower level: the line is known to exhibit redshifted absorption due to in-falling gas and/or blueshifted absorption due to inner disk winds, both of which are believed to be powered by accretion (Edwards et al. 2006). As summarized in Table 3, four sources showed Pa γ emission and six sources exhibited redshifted and/or blueshifted He I absorption. Nevertheless, these lines are

weak compared to those seen in accreting CTTS (e.g., Edwards et al. 2006). This implies low accretion rates, although edge-on inclinations may also affect line profiles. Some dippers (EPIC 204757338, 204932990, 205519771) show redshifted and/or blueshifted He I absorption but lack Pa γ emission and are classified as WTTS.

We checked for [O I] forbidden emission at 5577 Å and 6300 Å in our SNIFS and ESI spectra. These lines are hallmarks of UV-photodissociated OH and H₂O gas in the surface layers of inner- to mid-disk regions around T Tauri-like stars (e.g., Gorti et al. 2011). Only EPIC 204932990 showed weak emission in both lines with equivalent widths of ~ 1 –2 Å. To check whether this emission could be due to airglow features from the Earth’s atmosphere known to occur at the same wavelengths, we searched our collection of $\sim 2,000$ SNIFS spectra of late-type stars from the Lépine & Gaidos (2011) catalog. We found only 5 stars ($\sim 0.2\%$) with equivalent widths ≥ 0.5 Å, making it very unlikely that the detected [O I] emission from EPIC 204932990 was due to airglow. No dipper showed Ca II IRT emission in its optical spectrum.

4.3. Disk Dust & Gas Masses

Sub-mm continuum emission can be used to estimate the total mass of dust in a disk. Because emission at sub-mm wavelengths from protoplanetary disks is typically optically thin, the continuum flux (F_{ν}) is directly related to the total dust mass (M_{dust}) as in Hildebrand (1983):

$$M_{\text{dust}} \approx \frac{F_{\nu} d^2}{\kappa_{\nu} B_{\nu}(T_{\text{dust}})}, \quad (2)$$

where B_{ν} is the Planck function and we assume a characteristic dust temperature of $T_{\text{dust}} = 20$ K, which is the median found for Taurus disks (Andrews & Williams 2005). The dust grain opacity, κ_{ν} , is taken as 10 cm^2

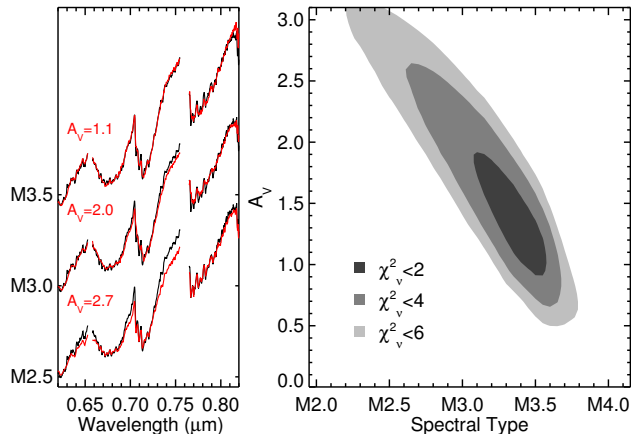


Figure 5. Example of our extinction fitting technique. Left: SNIFS spectrum of EPIC 204932990 (black lines) compared to de-reddened template spectra from TWA (red lines); regions with strong emission or telluric absorption features are masked out. Right: The χ^2 surface describing an expanded range of the fits shown on the left.

g^{-1} at 1000 GHz with an opacity power-law index $\beta = 1$ (Beckwith et al. 1990). We assumed distances, d , of 140 pc for Upper Sco and 120 pc for ρ Oph members.

We detected only three dippers with the SMA at $>3\sigma$ significance: EPIC 203937317 at 10.3 ± 1.8 mJy, EPIC 204630363 at 28.0 ± 1.5 mJy, and EPIC 204757338 at 12.5 ± 1.4 mJy. These sub-mm fluxes translate to dust masses of ~ 4 , 15, and 7 M_{\oplus} , respectively. EPIC 205151387 was previously observed with the SMA by Cieza et al. (2008) and ALMA by Carpenter et al. (2014), who measured sub-mm fluxes corresponding to dust masses of 7–9 M_{\oplus} . The 3σ upper limits for the undetected dippers translate to dust masses of $\lesssim 3 M_{\oplus}$.

Disk gas masses can be estimated from CO line emission. Converting integrated ^{12}CO 2–1 line emission into total gas mass is complicated by the optical thickness of the line. Instead, our best constraints come from comparing CO isotopologue line emission to the grid of models in Williams & Best (2014), who used these optically thin lines to predict total disk gas masses to within a factor of three. We did not detect ^{13}CO or C^{18}O isotopologue line emission in our SMA data for any of the dippers at ~ 40 mJy km s $^{-1}$ sensitivity. The 3σ upper limits only correspond to model disks from Williams & Best (2014) with gas masses $< 1 M_{\text{Jup}}$ (see their Figure 6), providing a rough upper limit to the dipper disk gas masses.

4.4. Infrared Color Excess & Variability

For a typical young star hosting a disk, near-IR continuum emission arises from hot dust at <1 AU, while mid-IR emission probes warm dust in the inner tens of AU, and emission at far-IR and mm wavelengths traces cold dust in the outer disk regions. Thus IR excesses above the stellar photosphere can be used to infer radial disk structure and classify disk evolutionary states. “Full” protoplanetary disks have not experienced significant clearing and thus exhibit strong excesses across the IR regime, indicating optically thick material throughout the disk. “Transition” disks lack IR excesses at $\lesssim 10 \mu\text{m}$ but show strong excesses at longer wavelengths, reflecting inner cavities formed by disk clearing mechanisms (e.g.,

planet formation); a subset of “pre-transition” disks exhibit near-IR excesses, indicating some optically thick material close to the star, but also mid-IR dips in their SEDs, indicating gaps in their disk structure (see review in Espaillat et al. 2014). “Evolved” disks have steadily decreasing IR excesses with increasing wavelength, reflecting their evolution toward optically thin disks. “Debris” disks are composed of second-generation dust created from collisions among planetesimals; their dust masses are typically several orders of magnitude lower than those of full protoplanetary disks. Note that for cool M dwarf stars, near-IR emission from the disk is weak compared to that from the stellar photosphere, thus near-IR excess from the inner disk may not be readily apparent.

Luhman & Mamajek (2012) used IR color excesses relative to photospheric colors to classify disks in Upper Sco. They computed IR colors using *Spitzer* and WISE bands (which are dominated by dust emission) relative to K band (which is dominated by stellar photosphere emission) then subtracted the IR color expected for a stellar photosphere of the same spectral subtype, as determined by fitting the sequence of diskless stars in Upper Sco (see their Figure 1). We calculated analogous IR color excesses using the W3 and W4 bands relative to the 2MASS K_S band: $E(K_S - W3)$ and $E(K_S - W4)$, respectively. Figure 6 plots these values against each other, indicating the different disk classifications for members of Upper Sco by color (Figure 6 is comparable to the bottom panel of Figure 2 in Luhman & Mamajek 2012). We over-plotted our dipper sample to assign initial disk types: the dippers appear to include full, evolved, and possibly transition disks. No dipper in our sample is classified as a debris or diskless system.

Some dippers in our sample exhibited mid-IR variability between their two W1/W2 observation epochs, taken six months apart. Four dippers (EPIC 203410665, 203895983, 204137184, 204630363) exhibited ~ 10 –15% increases/decreases in W1 and W2 flux over this time period, while the remaining dippers showed $\lesssim 5\%$ changes. (Pre-) transition disks are known to show so-called “see-saw” mid-IR variability, where flux at shorter wavelengths increases/decreases while flux at longer wavelengths decreases/increases. This has been interpreted as changes in the height of the inner disk wall (e.g., Espaillat et al. 2011; Flaherty et al. 2012) and could be related to the dipper phenomenon if the dimming events originate in the inner disk. Some young (~ 10 –200 Myr) debris disks have also shown mid-IR variability on similar scales, interpreted as debris from planetesimal collisions (e.g., Melis et al. 2012). For example, the young solar analog ID8 showed a $\sim 50\%$ brightening followed by a general decay over ~ 1 year at both IRAC-1 and IRAC-2 bands (Meng et al. 2014). This could also be related to the dipper phenomenon, if the dimming events originate from clumps of debris in the disk.

4.5. Spectral Energy Distributions

SEDs for each dipper are shown in Figure 7 with photometry given in Table 6. We derived the optical magnitudes from our SNIFS spectra by integrating over the revised filter profiles from Mann & von Braun (2015). We compared each SED to the expected stellar photosphere emission by plotting them against PHOENIX/CFIST

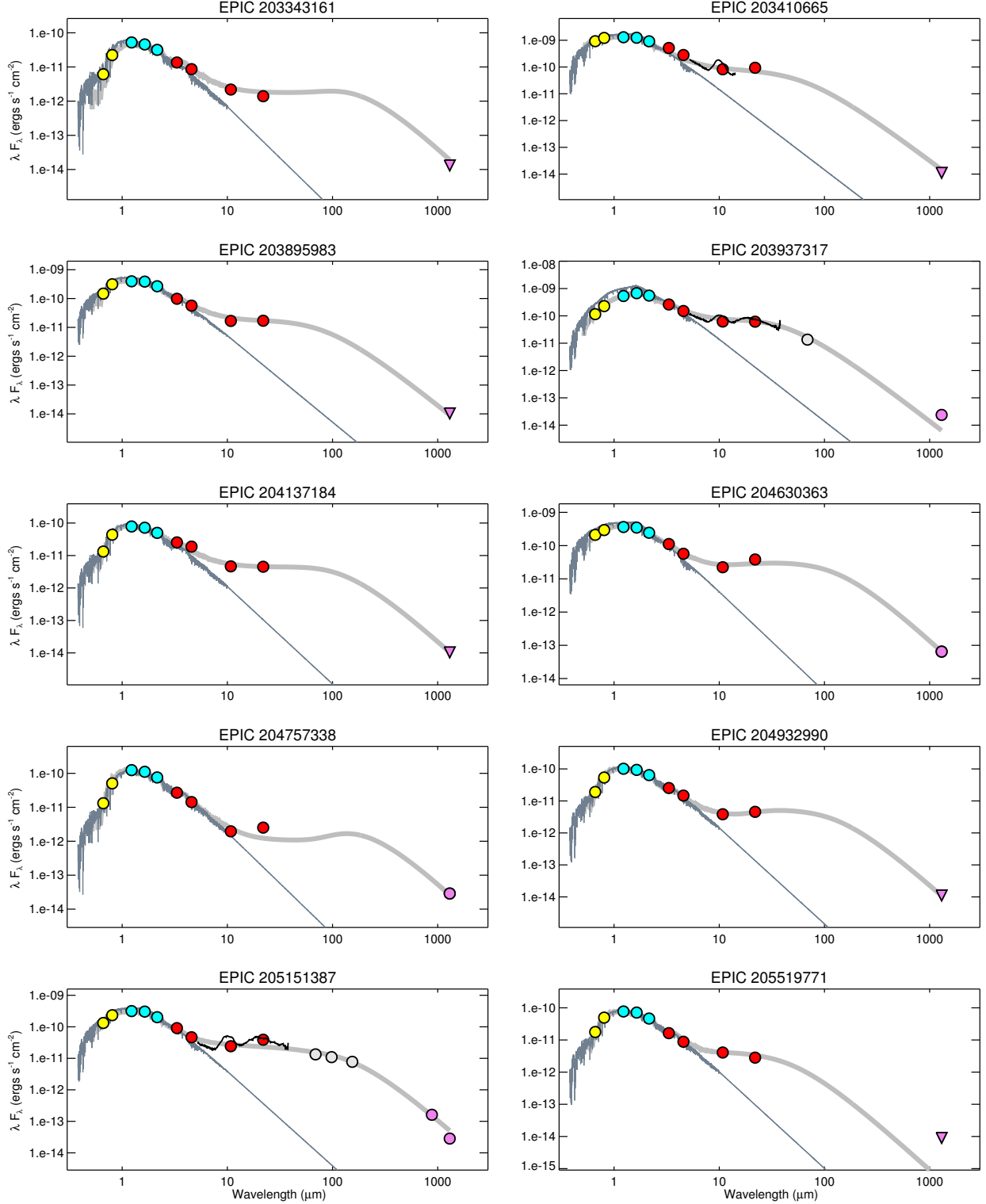


Figure 7. SEDs of the ten dippers in our sample. Yellow points are R_C and I_C magnitudes derived from our SNIFS spectra, blue points are 2MASS JHK_S magnitudes, and red points are WISE 1–4 magnitudes. *Spitzer* IRS spectra (black lines) and *Herschel*/PACS magnitudes (gray points) are shown when available. Purple points are SMA 1.3 mm fluxes, where downward triangles indicate upper limits (Section 3.2 & 3.4). EPIC 205151387 also has an ALMA observation at 880 μm from [Carpenter et al. \(2014\)](#) (Section 3.4). Errors on the photometry points are all smaller than the symbols; the photometry is summarized in Table 6. Dark gray curves show the expected stellar photosphere emission, derived from extinction-corrected PHOENIX/CFIST model spectra using the T_{eff} and A_V values in Table 1 and the reddening law from [Cardelli et al. \(1989\)](#). Thicker light gray curves show the SED models discussed in Section 4.5 with best-fit parameters given in Table 4.

Table 4
SED Modeling Results

EPIC	T_{eff} (K)	i_c (deg)	A_V (mag)	a_{in} (R_*)	a_{out} (R_*)	$\log(\varepsilon)$
203343161	3050	1	1.3	9.6	1.5E+04	0.0
203410665	4660	15	1.0	1.1	7.4E+02	-1.3
203895983	3560	10	0.4	1.4	1.1E+03	-1.1
203937317	4200	4	3.0	1.0	7.2E+02	-0.6
204137184	3280	6	0.8	1.0	3.4E+03	-0.8
204630363	4160	17	1.1	24.4	3.3E+03	-1.2
204757338	2890	2	0.1	2.9	9.9E+03	-1.0
204932990	3280	55	0.7	38.8	1.2E+03	-1.6
205151387	3860	90	0.9	3.8	1.8E+03	-1.8
205519771	3300	52	0.5	8.1	1.7E+02	-1.6

(2015). EPIC 203895983 is a likely near-equal-mass binary ($\Delta K_S = 0.1$ mag) with $\rho = 0.298'' \pm 0.001''$ (~ 40 AU) and $\text{PA} = 250^\circ 41 \pm 0^\circ 21$. EPIC 204932990 has a very faint candidate companion ($\Delta K_S = 5.6$ mag) with $\rho = 1.080'' \pm 0.002''$ (~ 150 AU) and $\text{PA} = 232^\circ 14 \pm 0^\circ 07$. Errors were derived following Bowler et al. 2015. Additional observations (e.g., multi-epoch astrometry) are needed to confirm that these candidate companions are not background objects, however the probability of unassociated objects within $1''$ should be low as no other stars were seen in the $5'' \times 5''$ NIRC2 images. We also searched our ESI spectra for spectroscopic binaries; however, autocorrelations of the spectra after applying a high-pass filter showed no evidence for spectroscopic binaries for the seven dippers with ESI data.

For the candidate close binaries shown in Figure 8, it is plausible that both components host disks given their similar ages. Thus the IR excess could be coming from one or both of the companions, as the WISE resolution is much larger than the binary separations. However, if the dipper phenomenon only occurs for disks seen nearly edge on, then it would be unlikely for both stars to contribute to the dipping events: Hale (1994) showed that coplanarity between the equatorial and orbital planes begins to diverge for binary systems with separations > 10 AU (see their Figure 2). Moreover, our periodograms (Section 2.2) only show one distinct period.

5. TRENDS AND CORRELATIONS

5.1. Disk Structure

The dippers in our sample appear to host moderately evolved protoplanetary disks. For example, the dippers detected in the 1.3 mm continuum have dust masses of $\sim 4\text{--}15 M_\oplus$ (Table 3), which are an order of magnitude lower than typical protoplanetary disks at $\sim 10^2 M_\oplus$, yet an order of magnitude higher than typical debris disks at $\sim 10^{-1} M_\oplus$ (see Figure 3 in Wyatt 2008). Their upper limits on disk gas mass of $\sim 1 M_{\text{Jup}}$ also imply gas-to-dust ratios of $< 20\text{--}75$, which are lower than the canonical ISM value of 100 (Bohlin et al. 1978); however, these gas-to-dust ratios are likely uncertain by at least a factor of three. Still, the weak accretion signatures exhibited by most of the dippers in our sample suggest accretion states between those of CTTS and WTTS (Section 4.2; Table 3): $\text{H}\alpha$ emission was typically consistent with stellar youth rather than ongoing accretion; only one dipper showed evidence of forbidden [O I] emission; no dippers exhibited Ca II IRT emission; and the near-IR accretion signatures were weak compared to CTTS.

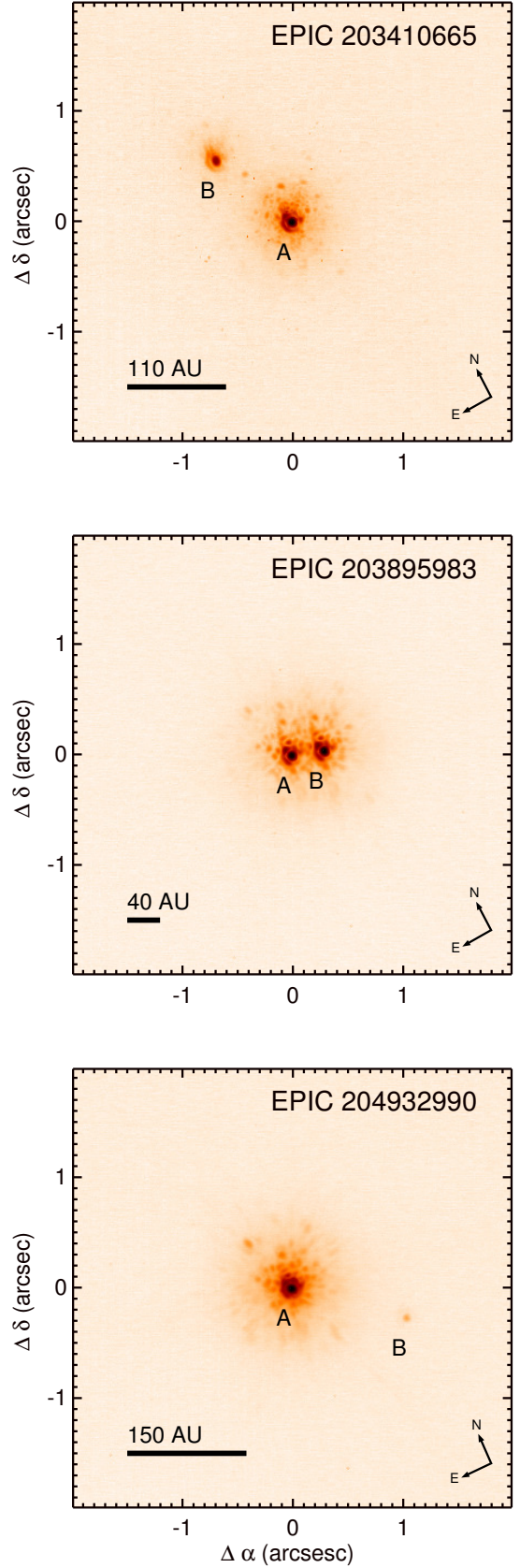


Figure 8. NIRC2 AO images for the dippers in our sample with candidate companions. Primary (A) and secondary (B) components are labeled. Scalebars showing separations in AU are plotted for reference. North is up and East is left.

Low disk masses are somewhat expected for the relatively old age of Upper Sco (~ 10 Myr). [Andrews et al. \(2013\)](#) showed that disk masses in Upper Sco are on average lower than those in the younger Taurus region (~ 2 Myr) by $\sim 2.5\sigma$. Still, such low disk masses could be due in part to close companions. [Harris et al. \(2012\)](#) showed that mm luminosities of Taurus disks decrease by $\sim 5\times$ for binary separations of ~ 30 – 300 AU and by another $\sim 5\times$ for separations of < 30 AU. They interpreted this trend as being due to the tidal truncation of circumstellar disks in close stellar systems. The three dippers whose AO images revealed candidate companions in the ~ 30 – 300 AU range have some of the lowest disk masses in our sample, yet the two other dippers with AO images have similarly low disk masses and appear to be single stars. Thus although close companions may have influenced the disks around some of our dippers, they cannot explain the low disk masses seen across our sample.

Moreover, half of our dippers show mid-IR dips in their SEDs, indicating they may host (pre-) transition disks that have developed gaps/holes in their disk structures (Table 3). Our SED modeling (Section 4.5) also suggests that EPIC 204932990 and 204630363 have inner holes, but that the remaining dippers have inner disks extending to within a few stellar radii. Note that our SED models do not account for pre-transition disks, which have gaps between inner/outer disk regions, as evident by the inability of the SED models to reproduce the mid-IR dips in some cases (e.g., EPIC 205151387; Figure 7).

5.2. Correlations between disk and dip properties

Figure 9 shows $E(K_S - W2)$ vs. D_{dip} for all the late-type members of Upper Sco listed in [Luhman & Mamajek \(2012\)](#) that were also observed during K2/C2. D_{dip} is a measure of dip depth, while $E(K_S - W2)$ is an indicator of dusty material in the inner disk. D_{dip} was derived from the K2/C2 light curves as in Section 2.2, and $E(K_S - W2)$ was calculated from the magnitudes listed in [Luhman & Mamajek \(2012\)](#) as in Section 4.4. The red line shows the D_{dip} cutoff that we used for identifying dippers (Section 2.2); points above the cutoff include our dipper sample (black diamonds), additional dippers flagged by subsequent visual inspection of the K2/C2 light curves (orange points; see Table 5), and sources that did not meet our other dipper criteria and/or had corrupted light curves (gray points).

Figure 9 shows a positive correlation between D_{dip} and $E(K_S - W2)$ for the dippers (a Spearman rank test gives $\rho = 0.51$ with a p-value of 0.01). The W2 band at $4.6 \mu\text{m}$ corresponds to a blackbody temperature of ~ 600 K, which is roughly the temperature of dust grains orbiting at a few stellar radii around an early-M dwarf. Our SED modeling shows that most of our dippers likely have inner disks extending down to these small distances (Table 4). We therefore interpret this correlation as evidence that the dips are related to inner disk material occulting the star. Indeed, this correlation with dip depth does *not* hold for W3 and W4 bands, which correspond to cooler material further out in the disk. Figure 9 also shows that sources can have inner disks but not exhibit dips, which may be a geometric effect: if the dipper phenomenon requires disks with edge-on geometries, then stars with large W2 excess may not exhibit dips.

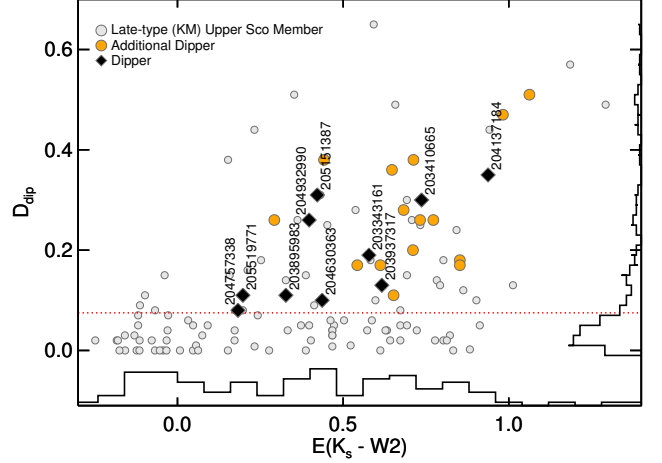


Figure 9. W2 excess versus dip depth for late-type Upper Sco members in [Luhman & Mamajek \(2012\)](#) that were observed during K2/C2 (gray points). The red line shows our D_{dip} cutoff used for dipper identification (Section 2.2). Our ten dippers (black diamonds; Table 1) and the additional dippers (orange points; Table 5) are over-plotted. Histograms show the distributions of points in the D_{dip} – $E(K_S - W2)$ plane projected onto the D_{dip} and $E(K_S - W2)$ axes, respectively.

Table 5
Additional Dippers[†]

EPIC	2MASS
203429083	15570350-2610081
203824153	16285407-2447442
203843911	16262367-2443138
203850058	16270659-2441488
203862309	16274270-2438506
203969672	16270907-2412007
203995761	16281673-2405142
204107757	15560104-2338081
204211116	16214199-2313432
204329690	16220194-2245410
204449274	16222160-2217307
204489514	16030161-2207523
204530046	16105011-2157481
204864076	16035767-2031055
205068630	16111095-1933320

[†] These targets were found by subsequent visual inspection of the K2/C2 light curves and will be the subject of future follow-up observations.

5.3. Correlations with Stellar Properties

The ten dippers in our sample are all late-K or M dwarf stars (Table 3). This is unlikely to be the result of the community-based selection of the K2/C2 sample: of the $\sim 13,000$ long-cadence targets, $\sim 55\%$ have $K_p - J > 1.8$, which is approximately the late-K and M dwarf regime. Thus the probability of picking ten random stars and having them all be late-K or M dwarfs is only $\sim 0.1\%$. Moreover, $\sim 80\%$ of the early-type and $\sim 95\%$ of the late-type members of Upper Sco (based on the samples from [Luhman & Mamajek 2012](#) and [Rizzuto et al. 2015](#)) were observed in K2/C2, making it unlikely that all the dippers would be cool dwarfs by chance.

Rather, this bias toward late-type stars could result from the dipper phenomenon being related to circumstellar material, as disks persist significantly longer around

low-mass stars compared to high-mass stars (e.g. Ribas et al. 2015). At the ~ 10 Myr age of Upper Sco, the disks around early-type stars have experienced significant clearing: for example, Luhman & Mamajek (2012) found that the fraction of Upper Sco members with inner primordial disks is $\lesssim 10\%$ for early-type (BAFG) stars, but increases to $\sim 25\%$ for late-type (M5–L0) stars.

Furthermore, the bias toward late-type stars could imply that the dips are related to circumstellar material located at a specific distance from the star where a given temperature is reached. The lower stellar luminosities of M dwarfs mean that such temperature-dependent disk components exist closer to the star, where transiting orbits are more likely. Indeed, we find a positive correlation between P_{rot} and T_{eff} in our dipper sample: the cooler systems have $P_{\text{rot}} \sim 2\text{--}3$ days, while the hottest systems have $P_{\text{rot}} \sim 4\text{--}10$ days (a Spearman rank test gives $\rho = 0.77$ with a p-value of 0.01). These inferred rotation periods correspond to dust grains at $\sim 400\text{--}600$ K, which is where the W2 band peaks and we observe a correlation with dip depth (Figure 9), providing further support that the dipper phenomenon arises from the inner disk.

6. PROPOSED MECHANISMS

Figure 3 shows representative examples of the three main dip types seen in the K2/C2 light curves (symmetric, trailing tail, complex). We compare these dips to model planetary transits in order to highlight their differences. Although the symmetric dips have similar shapes to transits of giant planets with short orbital periods around small stars, their dip durations and depths are both too large to be planetary. Similarly, although the dips with trailing tails resemble occultations by exocomets (e.g., Lecavelier Des Etangs et al. 1999), their depths are several orders of magnitude deeper than expected. Below we discuss three possible mechanisms to explain the dipper phenomenon, all of which are related to dusty circumstellar material passing through our LOS to the star. No single mechanism can explain our entire dipper sample alone, yet collectively these mechanisms can account for the range of observed dipper properties; given the diversity of the dipper light curves, it may be that the dippers are produced by different mechanisms that correspond to different stages of disk evolution.

6.1. Occulting Inner Disk Warps

Occultations of a star by non-axisymmetric structures in the inner disk could explain the dipper phenomenon, if the disk is seen nearly edge on. Given the amplitudes of the observed dips, the occulting material must have a vertical scale comparable to the size of the star. Assuming the material is orbiting with a Keplerian period equal to the inferred stellar rotation period, P_{rot} , (i.e., near the co-rotation radius) the co-inclination must be $\lesssim 27 R_* P_{\text{rot}}^{-2/3} M_*^{-1/3}$ degrees for an occultation to occur, where R_* and M_* are in solar units and P_{rot} is in days. This equates to a maximum co-inclination of ~ 9 degrees for an M dwarf of $R_* = 0.5 R_\odot$ and $M_* = 0.5 M_\odot$ with $P_{\text{rot}} = 2.5$ days. This means that a favorable geometry for occultation will occur $\sim 16\%$ of the time. However, of the ~ 135 Upper Sco members with “full” disks that were observed during K2/C2 (based on samples from Luhman & Mamajek 2012 and Rizzuto et al. 2015), we found only ~ 10 ($\sim 7\%$) to be dippers.

One previously proposed explanation for the quasi-periodic and aperiodic dippers seen in the CTTS populations of young star-forming regions is occultation by an inner disk warp near the co-rotation radius, where the warp in the accretion disk arises from dynamical interactions with an inclined magnetosphere. In particular, McGinnis et al. (2015) explained the quasi-periodic and aperiodic dippers in the young (~ 3 Myr) NGC 2264 star-forming region as resulting from stable and unstable accretion regimes, respectively. In stable accretion regimes, the disk warp occults the star with each rotation, causing regular dips in the light curve (Bouvier et al. 1999). In unstable accretion regimes (Kurosawa & Romanova 2013), stochastic occultations of the star occur when dust is lifted above the mid-plane near the base of an accretion column that passes along our LOS. McGinnis et al. (2015) found that some dippers in NGC 2264 switched between stable and unstable accretion regimes, exhibiting both quasi-periodic and aperiodic dimming over timescales of a few years.

This occulting inner disk warp scenario can explain the quasi-periodic (e.g., EPIC 203937317) and aperiodic (e.g., EPIC 203410665) dippers in our sample with clear signs of accretion. To check the plausibility of this scenario for these dippers, we estimated the lower limit on the semi-major axis of a disk warp, a_{warp} , by setting the Keplerian orbital period to the dip duration, τ , as shorter rotation periods would imply that the warp extends further than the distance around its orbit. For a star with $M_* = 0.5 M_\odot$ and $R_* = 0.5 R_\odot$, this yields $a_{\text{warp}} > 4, 7, \text{ and } 11 R_*$ for $\tau = 0.5, 1.0, \text{ and } 2.0$ days. These values are comparable to the co-rotation radii derived for M dwarf dippers in NGC 2264 by McGinnis et al. (2015). This implies that disk warps extending around most of their orbit, and producing dips lasting up to a couple of days, will exist near the co-rotation radius, as expected for the occulting disk warp scenario. Nevertheless, this scenario is difficult to apply to the majority of dippers in our sample, which have weak accretion signatures and low disk masses.

6.2. Vortices Produced by the Rossby Wave Instability

Non-axisymmetric structures in the inner disk can also arise from the Rossby Wave Instability (RWI). Rossby waves grow around extrema in the inverse potential vorticity of disks (Lovelace & Romanova 2014) until they evolve into vortices (Meheut et al. 2010) that are stabilized by local pressure maxima (Paardekooper et al. 2010). These extrema can occur at the boundaries between turbulent, magnetized, and accreting regions in the disk and “dead zones” where flows are laminar and viscous transport is low (Varnière & Tagger 2006). Stauffer et al. (2015) proposed such vortices as a possible explanation for a subset of shallow, short-duration, periodic dippers in NGC 2264 (their “Model 3”) but concluded that the degree of obscuration was insufficient to explain the dip amplitudes. We revisit this problem below, arguing that RWI-driven vortices can explain the quasi-periodic dippers in our sample with weak accretion signatures (e.g., EPIC 204757338), which are not readily explained by the disk warp scenario (Section 6.1).

The 3D gas dynamical simulations by Richard et al. (2013) numerically confirm the analytic predictions of Lesur & Papaloizou (2009) and Lin (2013) that insta-

bilities initially proceed rapidly at large azimuthal wave numbers but eventually the vortices merge. Only vortices with large (>6) aspect ratios in the radial-azimuthal plane survive breakdown by the elliptic streamline instability, and these vortices will typically span 1 radian of the orbit (e.g., Figure 3 in Richard et al. 2013). These characteristics of RWI-generated vortices explain the single mode of the quasi-periodic dippers as well as their dip depths and durations, which require that the occulting structures cover a significant fraction of their orbit. Indeed, the quasi-periodic dips seen in our sample appear to span roughly 1 radian (Figure 4). Moreover, RWI-generated vortices extend vertically across the scale height H of the disk, so if the vortices are optically thick then they will obscure a fraction H/R_* of the star for nearly edge-on inclinations. For molecular hydrogen at ~ 600 K orbiting at 2.5 days around a $0.4 R_\odot$ star, this equates to a dip depth of 15%, which is characteristic of our dippers.

Such vortices can be efficient traps of dust, enhancing local surface densities by 1–2 orders of magnitude, although high dust concentrations may ultimately cause instabilities (Fu et al. 2014; Crnkovic-Rubsamen et al. 2015). The absolute dust mass required to produce a dip is modest: the minimum mass of dust with diameter d required to produce a dip of depth f is $2\pi f d R_*^2 \rho / 3$. For a 10% dip, this equates to $\sim 10^{17}$ kg (the mass of a small asteroid) for mm-sized grains with bulk densities of 3 g cm^{-3} . Assuming the vortex spans 1 radian of the orbit and has an aspect ratio of 6, this corresponds to a surface density of $10^{-2} \text{ g cm}^{-2}$. For comparison, a few M_\oplus of dust (based on our sub-mm observations) spread uniformly over a disk of radius ~ 5 AU (based on our SED fitting) has a surface density of order 1 g cm^{-2} .

However, one issue with the vortex model is explaining how the amplitude and shape of the dips change between one orbit and the next, as simulations generally show that the vortices persist for 100s of orbits. One possibility is that, because of disk flaring, we only observe dips produced by vortices extending well above the disk mid-plane, as the LOS to the disk mid-plane would be obscured by the disk itself (Richard et al. 2013). Another issue is that the locations of these vortices would correspond to disk temperatures of ~ 400 – 700 K, according to the inner disk edges derived from our SED modeling (Section 4.5). These temperatures are well below the vaporization temperature of Na and K (~ 1000 K) where the disk gas is expected to become electrically conducting and therefore susceptible to MRI-driven turbulence. Because of this transition, the distance in the disk where temperatures reach ~ 1000 K is a candidate dead zone edge where vortices are likely to emerge. Still, mid-plane temperatures may be higher if there is a heat source internal to the disk, such as the viscous dissipation of the gravitational energy of accretion.

6.3. Transiting Circumstellar Clumps

The disk warp (Section 6.1) and RWI-generated vortex (Section 6.2) scenarios are difficult to reconcile with the aperiodic dippers in our sample with weak accretion signatures and low disk masses (e.g., EPIC 205519771). This is because disk warps/vortices near the co-rotation radius should produce regular dimming events every few

days. Unstable accretion (related to disk warps; Section 6.1) is also unlikely for these sources and cannot explain the widely separated dimming events. Thus we explore an alternative scenario where aperiodic dimming events result from single transits of dusty clumps of circumstellar material embedded at a few AU in the disk.

Protoplanetary disks are expected to be intrinsically clumpy because the gas distribution can exhibit low-level inhomogeneities that are amplified in the dust distribution (e.g., Birnstiel et al. 2013). Concentrations of particles at specific locations can lead to streaming instabilities, which further enhance over-densities (e.g., Johansen et al. 2007). Gravitational instabilities may then form ~ 1000 km-diameter planetesimals (Johansen et al. 2011). These processes have received much attention in numerical studies, as they can explain how dust grows into planetesimals, overcoming barriers of radial drift and bouncing (e.g., Windmark et al. 2012). Here we consider whether these physical processes can result in disks that are intrinsically clumpy at the levels required to explain the aperiodic dippers.

If we assume an equatorial transit and a circular orbit with speed v_{orb} , we can write the clump transit time across the host star (i.e., the dip duration) as $\tau = 2(R_c + R_*)/v_{\text{orb}}$ days, where R_c is the radius of the clump in solar radii and R_* is the radius of the host star in solar radii. From this, and the assumption that the clump masses are much smaller than the stellar mass, we can derive an expression for the clump size:

$$R_c \approx 1.85 \tau \left(\frac{M_*}{a} \right)^{1/2} - R_*, \quad (3)$$

where a is the semi-major axis of the clump orbit in AU and M_* is the mass of the host star in solar masses. The correlation between clump size and semi-major axis from Equation 3 is illustrated in Figure 10 for a star of $M_* = 0.5 M_\odot$ and $R_* = 0.5 R_\odot$ with dip durations of $\tau = 0.5, 1.0$, and 2.0 days.

We can place lower limits on a and R_c based on the nature of the aperiodic dimming events, which imply that the clumps have orbital periods longer than the 80-day K2/C2 observing campaign. Combining the requirement of $P > 80$ days with Kepler’s third law implies a lower limit to the semi-major axis of the orbiting clumps:

$$a > (GM_*)^{1/3} \left(\frac{P = 80 \text{ days}}{2\pi} \right)^{2/3} \approx 0.36 M_*^{1/3} \text{ AU}. \quad (4)$$

A lower limit to R_c can be inferred from the minimum dip depth; using D_{dip} (Table 1) as a measure of minimum dip depth, the lower limit to the clump size is:

$$R_c > R_* D_{\text{dip}}^{1/2}. \quad (5)$$

We show these observational restrictions on a and R_c as gray regions in Figure 10. Note that the limit on R_c corresponds to a dip depth of $D_{\text{dip}} \sim 0.1$ and that these limits only apply to the aperiodic dippers in our sample.

Could the transiting clumps be related to the ~ 1000 km planetesimals formed via streaming instabilities? Although the planetesimals are too small to cause the observed dips, their Hill spheres may provide suitably sized obscuring clumps. Figure 10 shows the Hill spheres of planetesimals with diameters of $D_p = 100, 1000$, and

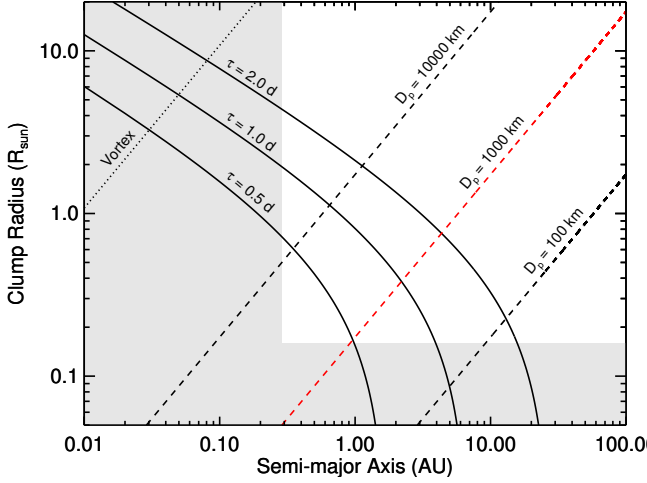


Figure 10. Relations between clump size (R_c) and semi-major axis (a). Dashed lines show the Hill spheres of planetesimals with different diameters, while solid lines show contours of constant dip durations; both assume an M dwarf host star with $M_* = 0.5 M_\odot$ and $R_* = 0.5 R_\odot$ (Equation 3). Gray regions show forbidden values of a and R_c for the transiting clump scenario from Equations (4) and (5), respectively. The dotted line shows an example of an RWI-generated vortex spanning 1 radian around its orbit for comparison.

10,000 km assuming bulk densities of 3 g cm^{-3} . The Hill spheres of $\sim 1000 \text{ km}$ planetesimals orbiting at a few AU can produce the observed dip durations within our constraints on R_c and a .

One concern with the transiting clump scenario is the required number of $\sim 1000 \text{ km}$ planetesimals, if the dips are indeed independent from each other. The total number of clumps in the disk that would eventually pass through our LOS can be estimated as:

$$N_{\text{los}} \approx \frac{P}{80 \text{ days}} N_{\text{obs}} \simeq 4.5 a^{3/2} M_*^{-1/2} N_{\text{obs}}, \quad (6)$$

where N_{obs} is the number of dips seen along our LOS during the 80-day K2/C2 observing period. To get the total number of clumps in the disk, N_{los} must also be corrected for the fact that the LOS only intercepts R_*/H of the disk for edge-on inclinations, where H is the scale height of the clump distribution in solar radii. Taking the disk scale height as $H/a \sim 0.1$, and using $N_{\text{obs}} \sim 15$ (from N_{dip} in Table 1), the total number of clumps in the disk is:

$$N_{\text{tot}} \approx 1000 a^{5/2} M_*^{-1/2} R_*^{-1}. \quad (7)$$

This predicts that $\sim 16,000$ clumps exist at $\sim 2 \text{ AU}$ for a star with $M_* = 0.5 M_\odot$ and $R_* = 0.5 R_\odot$. This corresponds to $\sim 4 M_\oplus$ or $\sim 20\%$ of the minimum mass solar nebula at that disk radius, which is fairly reasonable when compared to planetesimal formation models that predict $\sim 10\%$ of disk mass is converted into planetesimals (e.g., Johansen et al. 2011). Moreover it is implicit in many terrestrial planet formation models that much of the solid mass in protoplanetary disks is in planetesimals that eventually coalesces to form terrestrial planets.

Another concern is that the streaming instability operates in regions with high dust-to-gas ratios, namely the disk mid-plane where settling increases the dust density. This conflicts with the low extinction toward the dippers

(Table 1), which implies that we are not seeing their disks completely edge-on, and thus the transiting clumps are located above the disk mid-plane. However, high dust-to-gas ratios can also arise from gas depletion, thus the necessary environment for streaming instability may be possible at high altitudes if the gas has been sufficiently dispersed. Note that low disk densities are not necessarily an impediment to this process; for example, Carrera et al. (2015) found that the region in which planetesimal formation occurs should just move closer to the star as the disk dissipates. It could also be that planetesimals are scattered to high altitudes due to interactions with nearby planets (e.g., Krijt & Dominik 2011) or that the disk is warped due to an embedded inclined protoplanet (e.g., Facchini et al. 2014).

7. SUMMARY & FUTURE WORK

We have presented a sample of ten young ($\lesssim 10 \text{ Myr}$) late-K and M dwarf “dipper” stars located in the Upper Sco and $\rho \text{ Oph}$ star-forming regions. These sources were identified by dimming events in their K2/C2 light curves, which exhibited ~ 10 – 20 dips in flux over the 80-day observing campaign with typical durations of ~ 0.5 – 2 days and depths of up to $\sim 40\%$. We classified these dippers as either quasi-periodic or aperiodic: quasi-periodic refers to sources whose dips occur at periodic intervals but with varying shapes and depths; aperiodic refers to sources whose dips appear stochastically and with varying shapes and depths.

Our multi-wavelength follow-up observations revealed Li I absorption and $H\alpha$ emission consistent with stellar youth, but also rates of accretion spanning those expected from classical and weak-line T Tauri stars. All the dippers in our sample showed IR excesses consistent with protoplanetary disks, although our sub-mm observations implied low disk masses between those of typical protoplanetary and debris disks. SED modeling suggested that most dippers in our sample have inner disks extending to within a few stellar radii, although we could not rule out gaps in the disks, and two sources likely host disks with inner holes.

We found a positive correlation between dip depth and WISE-2 excess (but *not* WISE-3/4 excesses). We interpreted this as evidence that the dipper phenomenon arises from the inner disk near the co-rotation radius, where dust temperatures reach $\sim 600 \text{ K}$. However, this was difficult to reconcile with the aperiodic dippers in our sample because disk structures so close to the star should result in quasi-periodic dimming events. Although previous studies have reconciled aperiodic dippers among CTTS populations with unstable accretion regimes, the dippers presented in this work were mostly WTTS with low disk masses and weak accretion signatures.

We therefore explored several mechanisms that could be driving the dipper phenomenon: (1) inner disk warps near the co-rotation radius related to accretion; (2) RWI-driven vortices at the inner disk edge; and (3) transiting clumps of circumstellar material related to planetesimal formation at a few AU. For the quasi-periodic dippers, the transiting clump scenario is clearly ruled out as orbits near the co-rotation radius are needed to reproduce the regular dimming events with short periods. Rather, the quasi-periodic dippers are likely explained by occulting inner disk warps or RWI-driven vortices; however,

the requirement of accretion in the disk warp scenario may make RWI-driven vortices a more likely explanation for the quasi-periodic dippers in our sample with low disk masses and weak accretion signatures. Given the diversity of the dipper light curves, it could be that the dippers are produced by different mechanisms that correspond to different stages of disk evolution.

Additional follow-up observations may help distinguish among the proposed mechanisms driving the dipper phenomenon. High-resolution sub-mm imagery will reveal disk structures and geometries, while high-resolution spectra will better constrain stellar rotation periods and accretion states. Multi-wavelength, time-series photometry will also help determine dust grain properties. Moreover, follow-up observations of the 15 dippers found in subsequent visual inspection of the K2/C2 light curves will more than double our sample size. Such detailed analyses of dipper populations in young star-forming regions are important as they provide a unique opportunity to study inner disk structure and dynamics during key epochs of planet formation.

MCA and JPW were supported by NSF and NASA grants AST-1208911 and NNX15AC92G, respectively. EG was supported by an International Short Visit award from the Swiss National Science Foundation and by the Institute for Theory and Computation at Harvard CfA. MCW and GMK acknowledge European Union support through ERC grant 279973. TSB was supported through NASA grants ADAP12-0172 and ADAP14-0245. DML and TLJ thank Allan R. Schmitt for providing the light curve analysis software LcTools used in this study. We thank Fei Dai for checking light curves for planetary signals and Roberto Sanchis-Ojeda for providing us with photometry from his pipeline for many of our targets. This work was based on observations obtained with the Infrared Telescope Facility, which is operated by the University of Hawaii under Cooperative Agreement NNX-08AE38A with the NASA Science Mission Directorate's Planetary Astronomy Program.

REFERENCES

- Aldering, G., Adam, G., Antilogus, P., et al. 2002, in *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 4836, *Survey and Other Telescope Technologies and Discoveries*, ed. J. A. Tyson & S. Wolff, 61–72 [3.1](#)
- Alencar, S. H. P., Teixeira, P. S., Guimarães, M. M., et al. 2010, *A&A*, 519, A88 [1](#)
- Allard, F., Homeier, D., Freytag, B., et al. 2013, *Memorie della Societa Astronomica Italiana Supplementi*, 24, 128 [4.1](#)
- Andrews, S. M., Rosenfeld, K. A., Kraus, A. L., & Wilner, D. J. 2013, *ApJ*, 771, 129 [5.1](#)
- Andrews, S. M., & Williams, J. P. 2005, *ApJ*, 631, 1134 [4.3](#)
- . 2007, *ApJ*, 671, 1800 [1](#)
- Ardila, D., Martín, E., & Basri, G. 2000, *AJ*, 120, 479 [2.4](#)
- Aumer, M., & Binney, J. J. 2009, *MNRAS*, 397, 1286 [4.1](#)
- Baglin, A., Auvergne, M., Boissard, L., et al. 2006, in *COSPAR Meeting*, Vol. 36, 36th COSPAR Scientific Assembly, 3749 [1](#)
- Beckwith, S. V. W., Sargent, A. I., Chini, R. S., & Guesten, R. 1990, *AJ*, 99, 924 [4.3](#)
- Birnstiel, T., Dullemond, C. P., & Pinilla, P. 2013, *A&A*, 550, L8 [6.3](#)
- Bochanski, J. J., Hennawi, J. F., Simcoe, R. A., et al. 2009, *PASP*, 121, 1409 [3.1](#)
- Bohlin, R. C., Savage, B. D., & Drake, J. F. 1978, *ApJ*, 224, 132 [5.1](#)
- Borucki, W. J., Koch, D., Basri, G., et al. 2010, *Science*, 327, 977 [1](#)
- Bouvier, J., Chelli, A., Allain, S., et al. 1999, *A&A*, 349, 619 [6.1](#)
- Bowler, B. P., Liu, M. C., Shkolnik, E. L., & Tamura, M. 2015, *ApJS*, 216, 7 [4.6](#)
- Boyajian, T. S., von Braun, K., van Belle, G., et al. 2012, *ApJ*, 757, 112 [4.1](#)
- Caffau, E., Ludwig, H.-G., Steffen, M., Freytag, B., & Bonifacio, P. 2011, *Sol. Phys.*, 268, 255 [4.1](#)
- Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, 345, 245 [4.1](#), [4.5](#), [4.5](#), [7](#)
- Carpenter, J. M., Ricci, L., & Isella, A. 2014, *ApJ*, 787, 42 [2.4](#), [3.4](#), [4.3](#), [7](#), [6](#)
- Carrera, D., Johansen, A., & Davies, M. B. 2015, *A&A*, 579, A43 [6.3](#)
- Chiang, E. I., & Goldreich, P. 1997, *ApJ*, 490, 368 [4.5](#)
- Cieza, L. A., Swift, J. J., Mathews, G. S., & Williams, J. P. 2008, *ApJ*, 686, L115 [2.4](#), [3.2](#), [3.4](#), [4.3](#), [6](#)
- Claret, A. 2004, *A&A*, 428, 1001 [3](#)
- Cody, A. M., Stauffer, J., Baglin, A., et al. 2014, *AJ*, 147, 82 [1](#), [2.3](#)
- Crnkovic-Rubsamen, I., Zhu, Z., & Stone, J. M. 2015, *MNRAS*, 450, 4285 [6.2](#)
- Cushing, M. C., Vacca, W. D., & Rayner, J. T. 2004, *PASP*, 116, 362 [3.1](#)
- Dahm, S. E. 2008, *The Young Cluster and Star Forming Region NGC 2264* (Astronomical Society of the Pacific), 966 [1](#)
- de Zeeuw, P. T., Hoogerwerf, R., de Bruijne, J. H. J., Brown, A. G. A., & Blaauw, A. 1999, *AJ*, 117, 354 [1](#)
- Dotter, A., Chaboyer, B., Jevremović, D., et al. 2008, *ApJS*, 178, 89 [4.5](#)
- Ducourant, C., Teixeira, R., Galli, P. A. B., et al. 2014, *A&A*, 563, A121 [4.1](#)
- Edwards, S., Fischer, W., Hillenbrand, L., & Kwan, J. 2006, *ApJ*, 646, 319 [4.2](#)
- Elliott, P., Huélamo, N., Bouy, H., et al. 2015, *A&A*, 580, A88 [4.6](#)
- Espaillet, C., Furlan, E., D'Alessio, P., et al. 2011, *ApJ*, 728, 49 [4.4](#)
- Espaillet, C., Muzerolle, J., Najita, J., et al. 2014, in *Protostars and Planets VI* (University of Arizona Press), 497–520 [4.4](#)
- Facchini, S., Ricci, L., & Lodato, G. 2014, *MNRAS*, 442, 3700 [6.3](#)
- Fazio, G. G., Hora, J. L., Allen, L. E., et al. 2004, *ApJS*, 154, 10 [1](#)
- Flaherty, K. M., Muzerolle, J., Rieke, G., et al. 2012, *ApJ*, 748, 71 [4.4](#)
- Fu, W., Li, H., Lubow, S., Li, S., & Liang, E. 2014, *ApJ*, 795, L39 [6.2](#)
- Furlan, E., Watson, D. M., McClure, M. K., et al. 2009, *ApJ*, 703, 1964 [2.4](#)
- Gaidos, E., Mann, A. W., Lépine, S., et al. 2014, *MNRAS*, 443, 2561 [4.1](#)
- Gorti, U., Hollenbach, D., Najita, J., & Pascucci, I. 2011, *ApJ*, 735, 90 [4.2](#)
- Hale, A. 1994, *AJ*, 107, 306 [4.6](#)
- Harris, R. J., Andrews, S. M., Wilner, D. J., & Kraus, A. L. 2012, *ApJ*, 751, 115 [5.1](#)
- Harrison, T. E., Coughlin, J. L., Ule, N. M., & López-Morales, M. 2012, *AJ*, 143, 4 [2.1](#)
- Hartmann, L., & Kenyon, S. J. 1996, *ARA&A*, 34, 207 [1](#)
- Herbst, W., Eislöffel, J., Mundt, R., & Scholz, A. 2007, *Protostars and Planets V*, 297 [2.2](#)
- Hildebrand, R. H. 1983, *QJRS*, 24, 267 [4.3](#)
- Ho, P. T. P., Moran, J. M., & Lo, K. Y. 2004, *ApJ*, 616, L1 [3.2](#)
- Howell, S. B., Sobek, C., Haas, M., et al. 2014, *PASP*, 126, 398 [1](#)
- Huang, C. X., Penev, K., Hartman, J. D., et al. 2015, *ArXiv e-prints*, arXiv:1507.07578 [2.1](#)
- Johansen, A., Klahr, H., & Henning, T. 2011, *A&A*, 529, A62 [6.3](#), [6.3](#)
- Johansen, A., Oishi, J. S., Mac Low, M.-M., et al. 2007, *Nature*, 448, 1022 [6.3](#)
- Kinemuchi, K., Barclay, T., Fanelli, M., et al. 2012, *PASP*, 124, 963 [2.1](#)
- Krijt, S., & Dominik, C. 2011, *A&A*, 531, A80 [6.3](#)
- Kurosawa, R., & Romanova, M. M. 2013, *MNRAS*, 431, 2673 [6.1](#)

- Lantz, B., Aldering, G., Antilogus, P., et al. 2004, in Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series, Vol. 5249, Optical Design and Engineering, ed. L. Mazuray, P. J. Rogers, & R. Wartmann, 146–155 [3.1](#)
- Lecavelier Des Etangs, A., Vidal-Madjar, A., & Ferlet, R. 1999, *A&A*, 343, 916 [6](#)
- Lépine, S., & Gaidos, E. 2011, *AJ*, 142, 138 [4.2](#)
- Lépine, S., Hilton, E. J., Mann, A. W., et al. 2013, *AJ*, 145, 102 [3.1](#), [4.1](#), [4.2](#)
- Lesur, G., & Papaloizou, J. C. B. 2009, *A&A*, 498, 1 [6.2](#)
- Lin, M.-K. 2013, *ApJ*, 765, 84 [6.2](#)
- Lodieu, N. 2013, *MNRAS*, 431, 3222 [2.4](#)
- Loinard, L., Torres, R. M., Mioduszewski, A. J., & Rodríguez, L. F. 2008, *ApJ*, 675, L29 [1](#)
- Lovelace, R. V. E., & Romanova, M. M. 2014, *Fluid Dynamics Research*, 46, 041401 [6.2](#)
- Luhman, K. L., & Mamajek, E. E. 2012, *ApJ*, 758, 31 [2.4](#), [4.4](#), [6](#), [5.2](#), [9](#), [5.3](#), [6.1](#)
- Mann, A. W., Feiden, G. A., Gaidos, E., Boyajian, T., & von Braun, K. 2015, *ApJ*, 804, 64 [3.1](#)
- Mann, A. W., Gaidos, E., & Ansdell, M. 2013, *ApJ*, 779, 188 [4.1](#)
- Mann, A. W., Gaidos, E., Lépine, S., & Hilton, E. J. 2012, *ApJ*, 753, 90 [3.1](#)
- Mann, A. W., & von Braun, K. 2015, *PASP*, 127, 102 [4.5](#), [6](#)
- Markwardt, C. B. 2009, in Astronomical Society of the Pacific Conference Series, Vol. 411, Astronomical Data Analysis Software and Systems XVIII, ed. D. A. Bohlender, D. Durand, & P. Dowler, 251 [2.3](#)
- Mathews, G. S., Pinte, C., Duchêne, G., Williams, J. P., & Ménard, F. 2013, *A&A*, 558, A66 [2.4](#), [3.4](#)
- McClure, M. K., Furlan, E., Manoj, P., et al. 2010, *ApJS*, 188, 75 [2.4](#)
- McGinnis, P. T., Alencar, S. H. P., Guimarães, M. M., et al. 2015, *A&A*, 577, A11 [1](#), [2.3](#), [6.1](#)
- Meheut, H., Casse, F., Varniere, P., & Tagger, M. 2010, *A&A*, 516, A31 [6.2](#)
- Melis, C., Zuckerman, B., Rhee, J. H., et al. 2012, *Nature*, 487, 74 [4.4](#)
- Meng, H. Y. A., Su, K. Y. L., Rieke, G. H., et al. 2014, *Science*, 345, 1032 [4.4](#)
- Morales-Calderón, M., Stauffer, J. R., Hillenbrand, L. A., et al. 2011, *ApJ*, 733, 50 [1](#)
- Paardekooper, S.-J., Lesur, G., & Papaloizou, J. C. B. 2010, *ApJ*, 725, 146 [6.2](#)
- Pecaut, M. J., Mamajek, E. E., & Bubar, E. J. 2012, *ApJ*, 746, 154 [1](#)
- Prochaska, J. X., Gawiser, E., Wolfe, A. M., Cooke, J., & Gelino, D. 2003, *ApJS*, 147, 227 [3.1](#)
- Rayner, J. T., Toomey, D. W., Onaka, P. M., et al. 2003, *PASP*, 115, 362 [3.1](#)
- Reggiani, M., Robberto, M., Da Rio, N., et al. 2011, *A&A*, 534, A83 [1](#)
- Ribas, Á., Bouy, H., & Merín, B. 2015, *A&A*, 576, A52 [5.3](#)
- Richard, S., Barge, P., & Le Dizès, S. 2013, *A&A*, 559, A30 [6.2](#)
- Rizzuto, A. C., Ireland, M. J., & Kraus, A. L. 2015, *MNRAS*, 448, 2737 [2.4](#), [4.1](#), [5.3](#), [6.1](#)
- Sicilia-Aguilar, A., Hartmann, L. W., Watson, D., et al. 2007, *ApJ*, 659, 1637 [4.5](#)
- Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, 131, 1163 [3.4](#)
- Stauffer, J., Cody, A. M., McGinnis, P., et al. 2015, *AJ*, 149, 130 [1](#), [6.2](#)
- Still, M., & Barclay, T. 2012, PyKE: Reduction and analysis of Kepler Simple Aperture Photometry data, Astrophysics Source Code Library, ascl:1208.004 [2.1](#)
- Tukey, J. W. 1977, *Exploratory data analysis* (Addison-Wesley) [2.2](#)
- Vacca, W. D., Cushing, M. C., & Rayner, J. T. 2003, *PASP*, 115, 389 [3.1](#)
- van Dam, M. A., Bouchez, A. H., Le Mignant, D., et al. 2006, *PASP*, 118, 310 [3.3](#)
- Vanderburg, A., & Johnson, J. A. 2014, *PASP*, 126, 948 [2.1](#)
- Varnière, P., & Tagger, M. 2006, *A&A*, 446, L13 [6.2](#)
- Weinberger, A. J., Anglada-Escudé, G., & Boss, A. P. 2013, *ApJ*, 762, 118 [4.1](#)
- White, R. J., & Basri, G. 2003, *ApJ*, 582, 1109 [3](#), [4.2](#)
- Williams, J. P., & Best, W. M. J. 2014, *ApJ*, 788, 59 [4.3](#)
- Windmark, F., Birnstiel, T., Güttler, C., et al. 2012, *A&A*, 540, A73 [6.3](#)
- Wizinowich, P. L., Le Mignant, D., Bouchez, A. H., et al. 2006, *PASP*, 118, 297 [3.3](#)
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *AJ*, 140, 1868 [3.4](#)
- Wyatt, M. C. 2008, *ARA&A*, 46, 339 [5.1](#)
- Yelda, S., Lu, J. R., Ghez, A. M., et al. 2010, *ApJ*, 725, 331 [3.3](#)

Table 6
Photometry of K2/C2 Dippers

EPIC ID	R _c ^a (mag)	I _c ^a (mag)	2MASS-J (mag)	2MASS-H (mag)	2MASS-K _S (mag)	WISE-1 (mag)	WISE-2 (mag)	WISE-3 (mag)	WISE-4 (mag)	PACS-70 (mJy)	PACS-100 (mJy)	PACS-160 (mJy)	SMA ^b (mJy)
203343161	15.95±0.01	14.06±0.01	12.17±0.02	11.53±0.02	11.17±0.02	10.76±0.02	10.27±0.02	8.78±0.03	7.24±0.13	—	—	—	< 5.7
203410665	10.50±0.01	9.71±0.01	8.69±0.02	7.95±0.06	7.52±0.02	6.81±0.07	6.49±0.02	4.84±0.02	2.67±0.02	—	—	—	< 4.8
203895983	12.49±0.01	11.19±0.01	9.98±0.03	9.22±0.03	8.85±0.02	8.61±0.02	8.23±0.02	6.57±0.02	4.52±0.03	—	—	—	< 4.5
203937317	12.76±0.01	11.54±0.01	9.65±0.03	8.61±0.04	8.06±0.02	7.55±0.03	7.17±0.02	5.15±0.02	3.13±0.04	310±15	—	—	10.3±1.8
204137184	15.11±0.01	13.33±0.01	11.73±0.02	11.04±0.02	10.67±0.02	10.08±0.03	9.43±0.02	7.96±0.02	5.96±0.05	—	—	—	< 4.5
204630363	12.11±0.01	11.26±0.01	10.07±0.03	9.33±0.03	8.95±0.03	8.47±0.02	8.23±0.02	6.25±0.02	3.64±0.02	—	—	—	28.0±1.5
204757338	15.12±0.02	13.17±0.02	11.23±0.02	10.57±0.02	10.22±0.02	10.02±0.02	9.73±0.02	8.90±0.03	6.59±0.10	—	—	—	12.5±1.4
204932990	14.71±0.01	13.10±0.01	11.45±0.02	10.76±0.03	10.40±0.02	10.08±0.02	9.70±0.02	8.17±0.03	5.95±0.05	—	—	—	< 4.8
205151387	12.61±0.01	11.51±0.01	10.22±0.02	9.48±0.02	9.15±0.03	8.70±0.02	8.44±0.02	6.17±0.02	3.62±0.02	306±7	356±6	397±25	12.3±2.0
205519771	14.79±0.02	13.19±0.02	11.75±0.03	11.05±0.05	10.75±0.04	10.56±0.02	10.25±0.02	8.10±0.02	6.47±0.07	—	—	—	< 3.9

^a Cousins magnitudes estimated from our SNIFS optical spectra (Section 3.1) using the revised filter profiles and zero points from Mann & von Braun (2015).

^a Detected fluxes or 3 σ upper limits from our SMA 1.3 mm observations (Section 3.2), except for EPIC 205151387 whose SMA 1.3mm measurement comes from Cieza et al. (2008) and was also detected by ALMA at 880 μ m with 47.28±0.91 mJy (Carpenter et al. 2014).